

**exoALMA. XXI. The Morphology and Dynamics of Vertical Flows**

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Abstract

Vertical gas flows—such as winds and meridional circulations—are natural outcomes of protoplanetary disk processes and play a critical role in the earliest stages of planet formation. We analyze the vertical gas motions in 14 disks, as part of the exoALMA Large Program, focusing on the $^{12}\text{CO } J = 3-2$ and $^{13}\text{CO } J = 3-2$ emission lines. Using DISCMINER to model the Keplerian velocity field, we extract line-of-sight velocity residuals and measure the radial and vertical components of the gas motion. Vertical motions are detected in most disks. Two types of patterns emerge: (1) oscillatory up/down flows, likely linked to instabilities; and (2) transitions from downward to upward motions that we interpret as the bases of disk winds. In most cases, the velocity amplitudes are of a few tens of m s^{-1} . However, two disks, MWC758 and CQ Tau, show two spiral velocity features in their residual maps, red- and blueshifted, which we interpret as vertical velocities reaching up to $\sim 350 \text{ m s}^{-1}$ ($\sim 0.7c_s$), consistent with the gas motion in eccentric disks. Fast upward motions (up to 500 m s^{-1} ; $\sim 1.8c_s$) are also detected in the outer disk of MWC758. Synthetic observations from (magneto)hydrodynamic simulations validate the reliability of our method. Although strong molecular winds appear to be relatively rare in ^{12}CO and ^{13}CO , our study shows that, when traced by deep high-spectral-resolution line data, protoplanetary disks exhibit ubiquitous vertical flows. However, their overall velocity structure is highly complex, preventing the identification of a coherent dominant physical mechanism driving the vertical motions across all disks, thus requiring further theoretical investigation.

Unified Astronomy Thesaurus concepts: [Protoplanetary disks \(1300\)](#)

1. Introduction

The gas dynamics in a protoplanetary disk is dominated by Keplerian rotation around the central young star, and such a

velocity structure has been used to identify disks themselves (D. W. Koerner et al. 1993; S. Guilloteau & A. Dutrey 1994). Keplerian rotation has also routinely been used to obtain dynamical mass estimates for the central star and determine global disk properties (A. Dutrey et al. 1998; A. Isella et al. 2007; S. M. Andrews et al. 2024). In recent years, deviations from Keplerian rotation were proposed as a method to identify companions or planets still embedded in their host disks (C. Pinte et al. 2018; R. Teague et al. 2018), as a range of non-Keplerian motions, including vertical and radial components,

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can naturally emerge from planet–disk interactions (e.g., J. Bae et al. 2025). Once the Keplerian contribution is removed from the disk velocity field, the velocity residuals can be decomposed into three azimuthally symmetric orthogonal velocity components (azimuthal, vertical, and radial velocities). Azimuthal velocity residuals can, for example, result from pressure variations, locally (e.g., gaps) or globally (e.g., pressure drop-offs in outer disks and self-gravity; e.g., C. Longarini et al. 2025; J. Stadler et al. 2025).

Vertical motions in protoplanetary disks can be induced by various processes that result from the local combined effects of gravity, pressure, turbulence, and magnetic fields. For example, while the vertical shear instability (VSI) will result in clear upward and downward motions, the magnetorotational instability (MRI) will induce uncorrelated motions (M. Flock et al. 2015; M. Barraza-Alfaro et al. 2024, 2025). Winds, of magnetohydrodynamical (MHD) or photoevaporative (PE) nature, naturally arise, depending on the strength of the disk magnetic field and the level of stellar activity. MHD winds play a fundamental role in the evolution of protoplanetary disks, as they regulate angular momentum and allow material to accrete onto the star (G. Lesur et al. 2023; X. Hu et al. 2025a). PE winds in turn play an important role in the late-stage dispersal of disk material (R. Alexander et al. 2014). Both have dramatic impacts on the timescales and local conditions for planet formation and migration (X. Hu et al. 2025b; G. Wafflard-Fernandez & G. Lesur 2025). Upward and downward flows, such as meridional circulation flows, are expected to occur in planetary gaps (A. Morbidelli et al. 2014; E. Lega et al. 2024) and to deliver the material that will constitute the planetary atmospheres and define its chemical complexity (A. J. Cridland et al. 2025). All of these processes overall directly affect the disk thermal and chemical structure, dust mixing, grain growth, and settling and can lead to substructure formation (e.g., A. Riols et al. 2020).

Observational evidence for winds launched in regions close to the star has accumulated for decades at optical and infrared wavelengths (I. Pascucci et al. 2023, 2025). This is particularly clear in the forbidden lines that show a low-velocity blueshifted emission (e.g., I. Pascucci et al. 2020). Additional evidence for winds was found in the CO fundamental M -band emission line that shows broadened line wings and a low-velocity component (J. E. Bast et al. 2011; A. Banzatti et al. 2022). At millimeter wavelengths, large-scale outflows are also routinely found at the Class I/II stage (B. Tabone et al. 2017; M. Fernández-López et al. 2020; F. Bacciotti et al. 2025), sometimes with clear evidence of corotation with the disk (F. Louvet et al. 2018; A. de Valon et al. 2020; B. Tabone et al. 2020).

Evidence for disk winds in Class II objects is, on the other hand, rather sparse at these wavelengths, even though models incorporating MHD wind-driven accretion better match the observed disk properties than turbulent models (B. Tabone et al. 2025). Recent high-angular-and-spectral-resolution observations of disks with substructures have allowed for vertical motions to be mapped. In HD100546, S. Casassus et al. (2022) found a so-called “Doppler flip” (local sign change in velocity residuals) located in the dust continuum emission (see also B. J. Norfolk et al. 2022). Using multiple CO isotopologs, R. Teague et al. (2019), H. Yu et al. (2021), A. F. Izquierdo et al. (2023), and M. Galloway-Sprietsma et al. (2023) found strong upward flows (so-called “meridional

fountains”) in gas gaps located at distances larger than 100 au in two Herbig disks, HD163296 and HD169142, and the T Tauri disk AS209. Interestingly, R. Teague et al. (2021) and A. F. Izquierdo et al. (2023) found both upward and downward vertical motions of a few tens of m s^{-1} in MWC480 in the $^{12}\text{CO } J = 2-1$ line, while the other isotopologs tracing lower emitting heights did not show similar behaviors.

In this Letter, we expand such an analysis by considering the $^{12}\text{CO } J = 3-2$ and $^{13}\text{CO } J = 3-2$ emission-line observations for 14 out of the 15 disks in the exoALMA program (AA Tau being excluded from the analysis—see below). exoALMA constitutes an unprecedented dataset specifically designed for the detailed study of kinematics (R. Teague et al. 2025). The Letter is structured as follows. We first present the data that we use (Section 2) and the method employed to extract vertical motions (Section 3) that we benchmark in Section 4. In Section 5, we report the vertical motions measured in an axisymmetric analysis for 11 targets and those extracted along spiral spines for three targets. We discuss our findings in the discussion (Section 6) and conclude.

2. Data

We refer to R. A. Loomis et al. (2025) for the details of the calibration and imaging of the exoALMA observations. In this Letter, we focus on the $^{12}\text{CO } J = 3-2$ and $^{13}\text{CO } J = 3-2$ continuum-subtracted cubes. Similar to J. Stadler et al. (2025), we consider cubes with different angular resolutions, depending on the signal-to-noise ratio. We focus mainly on the so-called “Fiducial Images,” with beam size $0''.15$ and channel spacing of 100 m s^{-1} , and on the “High Surface Brightness Sensitivity Images” (beam size $0''.30$, channel spacing 100 m s^{-1} ; see R. Teague et al. 2025 for details). As we are mostly interested in following the large-scale features in this Letter, we favor a high signal-to-noise ratio at the cost of angular resolution when selecting the imaging setup. We report the cube angular resolution for all targets in Appendix A (Table A1). As shown in J. Stadler et al. (2025), using cubes at different angular resolutions from the exoALMA images does not significantly affect the overall velocity extraction. We exclude data within the central beam from the analysis.

We also consider three $^{12}\text{CO } J = 2-1$ data cubes for HD163296, MWC480, and AS209 from the MAPS Large Program (K. I. Öberg et al. 2021), with $0''.15$ angular resolution and 200 m s^{-1} channel spacing, for which kinematical analysis was previously published in A. F. Izquierdo et al. (2023). In addition, and only for the purpose of discussing the targets’ systemic velocities, we include the CS $J = 7-6$ cubes from exoALMA, with $0''.3$ and 200 m s^{-1} spatial and spectral resolutions, respectively, as well as the C^{18}O cubes from MAPS with similar properties as the ^{12}CO cubes. The full analysis of the CS $J = 7-6$ cubes will appear in a forthcoming publication from the collaboration.

3. Method

3.1. Velocity Extraction

The line-of-sight velocity can be expressed in terms of its three orthogonal velocity components— v_ϕ , v_r , and v_z —defined in the disk frame of reference:

$$v_{\text{los}} = v_{\phi, \text{proj}} + v_{r, \text{proj}} + v_{z, \text{proj}} + v_{\text{sys}}, \quad (1)$$

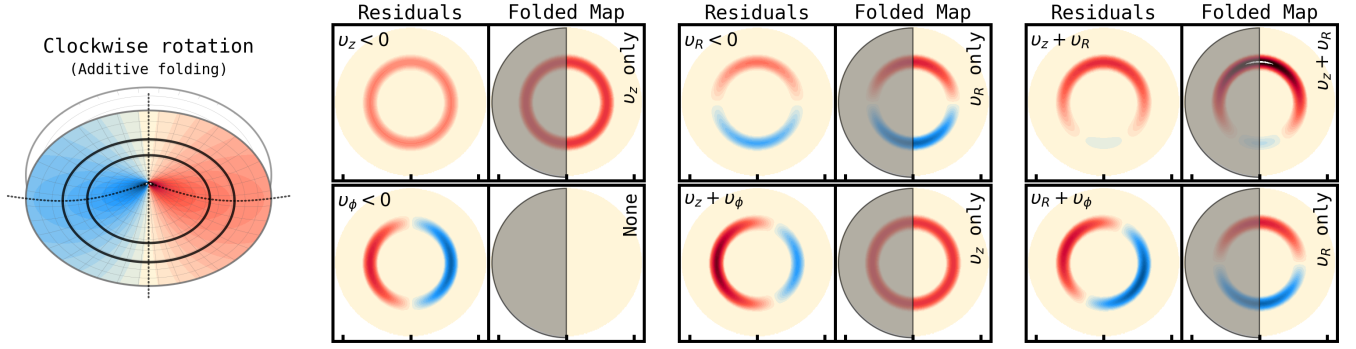


Figure 1. Sketch explaining the types of velocity residuals expected after the subtraction of a Keplerian disk model, in 2D residual maps (labeled “residuals”) and in folded residual maps (labeled “folded”). A more comprehensive sketch encompassing more cases can be found in (A. F. Izquierdo et al. 2026).

with

$$v_{\phi, \text{proj}} = v_{\phi} \cos(\phi) \sin(i) \cdot \text{sgn}_{\text{rot}}, \quad (2)$$

$$v_{r, \text{proj}} = -v_r \sin(\phi) \sin(i), \quad (3)$$

$$v_{z, \text{proj}} = -v_z \cos(i), \quad (4)$$

where v_{sys} is the systemic velocity, ϕ is the azimuthal angle of the disk measured counterclockwise from the redshifted major axis, $i \in [-90^\circ, 90^\circ]$ is the disk inclination, and sgn_{rot} denotes the sign of the disk rotation direction, with clockwise rotation defined as positive (see A. F. Izquierdo et al. 2025). Following these equations, an axisymmetric azimuthal (radial) velocity perturbation will appear in a 2D v_{los} map as two symmetric arcs with opposite-velocity signs relative to the projected minor (major) axis of the disk, while an axisymmetric vertical perturbation will manifest as a full annulus across all disk axes without changing signs. Figure 1 illustrates these three examples (v_z only, v_{ϕ} only, and v_r only).

We follow the procedure outlined in A. F. Izquierdo et al. (2025) and use the DISCMINER code, which fits a Keplerian disk model to the channel maps of a line intensity cube, assuming radial parametric forms for the peak intensity, line width, and emission surface height of the disk (A. F. Izquierdo et al. 2021). In the analysis, we consider only smooth disks with no warped regions. As in all exoALMA Letters, the geometrical parameters and the systemic velocity used in the ^{13}CO models are fixed to those derived from the ^{12}CO analysis. The best-fit models for all targets are presented in A. F. Izquierdo et al. (2025). At each location in the disk, the line-of-sight velocity is estimated by fitting the line profile with a Gaussian, a Bell, or a double-Bell function, depending on whether there is significant contribution from the back side of the disk to the observed emission. A Keplerian model is then subtracted from the data to produce a 2D residual v_{los} map, encoding velocity fluctuations in all three dimensions: δv_{ϕ} , v_r , and v_z . An analysis of the radial δv_{ϕ} profiles and their relations to gas pressure substructures can be found in J. Stadler et al. (2025).

Analyses of the velocity components have commonly been performed following a 1D axisymmetric extraction (e.g., R. Teague et al. 2019; H. Yu et al. 2021; M. Galloway-Sprietsma et al. 2023). Radial profiles of the azimuthal and vertical components are obtained by assuming that they are axisymmetric and orthogonal (for more details, see A. F. Izquierdo et al. 2023). Assuming the azimuthal symmetry of the azimuthal and radial velocity components, the vertical velocity component is simply the azimuthal average of the residual v_{los}

map divided by $-\sec(i)$ (A. F. Izquierdo et al. 2022). The radial velocities are obtained through an azimuthal average of all remaining velocity residuals that include v_r as well as the remaining nonaxisymmetric contributions from δv_{ϕ} and v_z . However, as noted in A. F. Izquierdo et al. (2023), when velocity perturbations strongly deviate from axisymmetry, such decomposition methods fail at properly retrieving the right amplitude for each velocity component. This motivates two different approaches for the axisymmetric and nonaxisymmetric residuals that we will explore in the following sections.

3.2. Residual Maps

2D residual maps. Among the 15 exoALMA targets, a large variety of line-of-sight velocity residuals are observed, as indicated in Figure 2, some showing significant nonaxisymmetric features (see also M. Fukagawa et al. 2026). Roughly speaking, there are three types of residual features: large spirals (as in MWC758; top left in Figure 2), ring-like features (as in J1604; bottom left), and patches/arcs (as in HD135344B; second row, left).

Folded residual maps. To quantitatively assess the presence and prominence of asymmetries and vertical velocity flows in our sources, we use folded residual maps, in which the line centroid velocity residuals from one half of the disk are added to those of the opposite half relative to the disk minor axis (A. F. Izquierdo et al. 2026). This procedure removes the typically dominant symmetric azimuthal velocity perturbations, such as those caused by pressure drop-offs in the outer disk or pressure perturbations due to gaps (J. Stadler et al. 2025), leaving only nonaxisymmetric motions and the axisymmetric component of v_z in the resulting map. Figure 1 illustrates the resulting folded maps for cases with only vertical, only azimuthal, only radial velocity perturbations, and combinations of those. In folded residual maps, any coherent vertical flow will appear as a single-colored semicircle.

We compute the folded residual maps for all of our targets and order them by the mean and standard deviation of their folded residual maps, as very simple metrics to estimate their dynamical state. We report the values for the ^{12}CO and ^{13}CO maps in Table A1 of Appendix A. We note that the targets that exhibit clear spiral features in their velocity residual maps (MWC758, CQ Tau, and SY Cha) show both the highest means and standard deviations of residual velocities. In contrast, J1604, PDS66, and V4046 Sgr yield the smallest mean and standard deviation values, indicative of uniform and

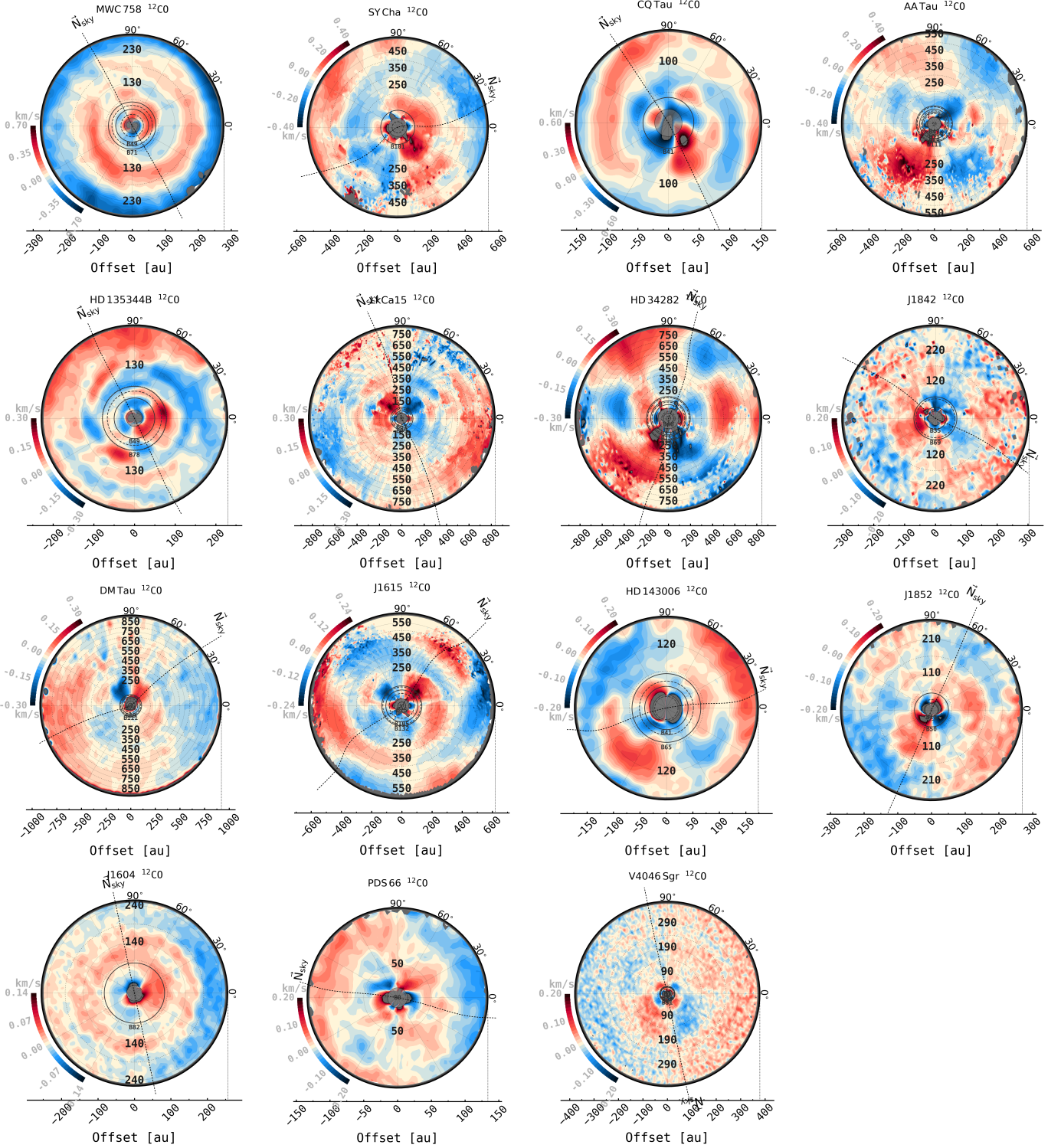


Figure 2. ^{12}CO line-of-sight residual maps for exoALMA targets. Note the different ranges of velocity amplitudes. Full and dashed circles give locations of continuum rings and gaps, respectively.

low-amplitude flows, suggesting that these sources are comparatively less dynamically active.

Samples for 2D and 1D analysis. In Section 5, we will present the velocity extraction for our sample. For a subset of three targets, for which clear velocity spirals are observed in the 2D residual maps, we trace coherent sub- and super-Keplerian substructures (filaments) using FILFINDER (E. W. Koch & E. W. Rosolowsky 2015). We then extract

the peak velocity residual along the spine of the filaments as well as the distribution of velocities across the spiral width. We consider a filament to trace vertical motions when the velocity does not change sign across either the minor or major axes, as azimuthal and radial velocity components would (see Equations (2) and (3); R. Teague et al. 2022). This applies to MWC758, CQ Tau, and SY Cha, though only marginally in the latter case. As explained previously, in the 2D and folded

residual maps, such a filament would cover a large azimuthal extent without changing velocity sign.

For all the remaining targets, except AA Tau (see below), we use the 1D axisymmetric extraction method, as in previous studies. We note that HD135344B shows a variety of arcs similar to short spirals, but since the velocity spirals change sign when crossing the major and minor axes, we do not consider those as coherent vertical motions. This target was explored in L. Wölfer et al. (2025), and a follow-up study will explore these motions in more detail in the context of planet-disk interactions (A. F. Izquierdo et al. 2026).

Surface height retrievals from the channel maps performed by M. Galloway-Sprietsma et al. (2025) show that the emission in the outer regions of SY Cha and AA Tau is diffuse and may originate from a distinct physical component or result from a pronounced pressure drop. We therefore exclude these outer regions (specifically, radii larger than 360 and 269 au, respectively). In the case of AA Tau, the remaining velocity residual map after removing the diffuse emission (see Figures 2 and C1) displays a characteristic pattern consistent with disk back-side contamination (A. F. Izquierdo et al. 2025). Given the combination of compact line emission and very high inclination, we exclude AA Tau from further analysis.

4. Benchmarking against Models

To understand our ability to recover vertical motions in dynamically perturbed disks, like those in our sample, we first test our retrieval methodology on synthetic cubes generated from hydrodynamical simulations for which we know the exact velocities. We consider four models. First, a planet-disk-interaction 3D hydrodynamical model, using the same surface density and temperature profile as in R. Teague et al. (2019) but considering a $4 M_{\text{Jup}}$ planet at 240 au. Then, the 3D (magneto)hydrodynamical simulations for: (1) the VSI; (2) the MRI; and (3) the gravitational instability (GI), computed with the PLUTO code (A. Mignone et al. 2012) and presented in M. Barraza-Alfaro et al. (2025). We first produce synthetic observations after radiative transfer with RADMC3D (C. P. Dullemond et al. 2012), with a 100 m s^{-1} channel spacing, then fit them with DISCMINER, using the same procedure as for our observations. We then fold the line-of-sight velocity residual maps, which are presented in the top panels of Figure 3. Just as in our observations, the folded maps include any nonaxisymmetric δv_ϕ , v_r , as well as v_z , for which any coherent vertical flow will appear as a single-colored semicircle.

In the panels below the folded maps in Figure 3, we show four groups of two panels for each simulation. In each top panel, we show the vertical motions retrieved by our method (the orange line, labeled “ v_z ”) and the ones that are encoded in the hydrodynamical simulation at a given Z/R that correspond to ~ 2.5 pressure scale heights (the blue line, labeled “ v_z model”). In each bottom panel, we present a radial profile of the standard deviation of the folded map, used as a measure of the nonaxisymmetric motions in our sources. We find that even in cases where the standard deviation of the folded map is high (e.g., in MRI and GI simulations), indicating a significant degree of asymmetry, the true model vertical velocities are consistently recovered within a few tens of m s^{-1} . This indicates that in the context of these four simulations, the

contributions from residual δv_ϕ and v_r do not add coherently to mimic the residual pattern characteristic of a vertical flow.

While these simulations might not be fully representative of the dynamical states of the exoALMA sample, we find that there is likely a floor value of about 10 m s^{-1} below which it is difficult to retrieve accurate velocities. Appendix B shows a comparison between the vertical velocities extracted from the VSI simulation at a given height $Z/R = 0.24$ with those extracted after postprocessing and beam convolution, for infinite and exoALMA-like angular resolution. We find that beam convolution naturally limits the amplitudes of the velocities measured in our observations (see also T. Hilder et al. 2025), implying that the retrieved values in the analysis presented in this Letter likely underestimate the true vertical motions occurring in the disks of our sample and that multiple layers around the $\tau = 1$ surface contribute to the measured velocity. Having benchmarked and validated our methodology on models, we apply the same approach to the exoALMA observations in the following sections.

5. Results

5.1. Systemic Velocities

To get insights into the amplitudes of the v_z motions involved, we first compute the maximum vertical motions in upward / downward directions for all targets considered in the 1D analysis. Since we exclude AA Tau and the spiral targets (MWC758, CQ Tau, and SY Cha), we are left with 11 targets for which we will analyze the azimuthally averaged radial profiles: HD34282, HD135344B, LkCa15, J1615, DM Tau, J1604, J1842, J1852, HD143006, PDS66, and V4046 Sgr. For comparison, we also include HD163296, AS209, and MWC480 from the MAPS Large Program within the same analysis pipeline.

Figure 4 presents the maximum upward and downward vertical velocities retrieved in our analysis (see also Table 1). Within the exoALMA sample, the peak upward velocities range from $4.0 \pm 8.8 \text{ m s}^{-1}$ (HD143006) to $37.8 \pm 10.4 \text{ m s}^{-1}$ (HD135344B), while the peak downward velocities span from $-4.5 \pm 2.6 \text{ m s}^{-1}$ (V4046 Sgr) to $-81.3 \pm 34.0 \text{ m s}^{-1}$ (HD34282). Here, we consider the error as the standard deviation in v_z along the azimuth at the radial location of the maximum v_z to provide an idea of the velocity dispersion along the azimuth. In comparison, the MAPS sources exhibit significantly stronger vertical motions, with AS209 displaying the highest upward velocity of $115.9 \pm 8.7 \text{ m s}^{-1}$ and downward velocity of $-86.0 \pm 9.6 \text{ m s}^{-1}$ (see also M. Galloway-Sprietsma et al. 2023; A. F. Izquierdo et al. 2023), which cannot be attributed to the different spectral resolutions (200 and 100 m s^{-1} for MAPS and exoALMA, respectively).

Interestingly, as is immediately clear from Figure 4, we find that the amplitudes of the maximum upward and downward motions are generally comparable. The systemic velocities derived from our model fitting therefore represent midpoints between positive and negative extremes, and this implies a degeneracy in the Keplerian fitting process, whereby the actual vertical motions in the disk may be partially interpreted as the bulk velocity of the source relative to the observer, hence introducing a systematic offset in the measured v_z .

To assess such a possible offset in the vertical velocity measurements, we examine the systemic velocities derived from another molecular line that traces a different vertical

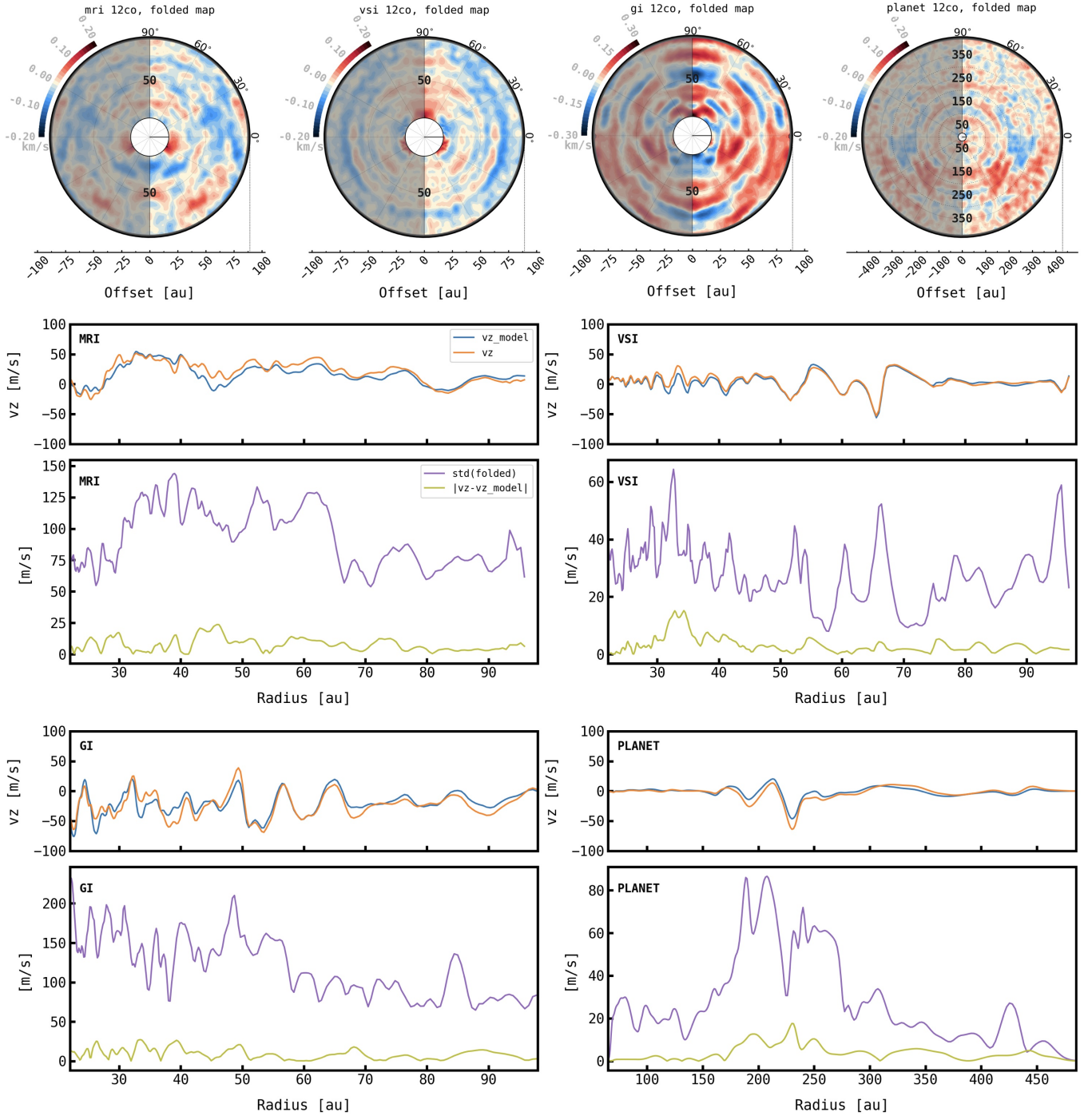


Figure 3. Benchmark of our v_z retrieval method with hydrodynamical models. At the top, we show folded residual maps from the MRI, VSI, GI, and planet–disk–interaction models. For each subset of two panels, we present the v_z from the model and the one retrieved with DISCMINER (top), as well as the standard deviation of the folded map and the difference between the model and measured v_z (bottom), all against radius.

layer of the disk, providing an independent measurement. Specifically, we model the CS $J = 7-6$ emission line, which, in most systems, predominantly probes regions closer to the midplane (see M. Galloway-Sprietsma et al. 2025). We proceed in a similar way as for the ^{12}CO cubes, fitting a Keplerian model to the channel maps using DISCMINER, with the systemic velocity as a free parameter. If the systemic velocities retrieved from both lines agree, the inferred vertical velocities likely reflect the true absolute gas motions.

However, any significant discrepancy between them would likely point toward an absolute offset affecting the v_z retrieval.

Table 1 summarizes the systemic velocities derived with DISCMINER for ^{12}CO and CS, also shown in Figure 4. Among the targets exhibiting clear vertical motions in our retrieval, we find systemic velocity differences between the two lines ranging from 0.5 m s^{-1} (J1852) to 20.5 m s^{-1} (HD135344B). For the MAPS sources, these differences, computed from the ^{12}CO and C^{18}O lines, are 10 m s^{-1} (HD163296 and MWC480)

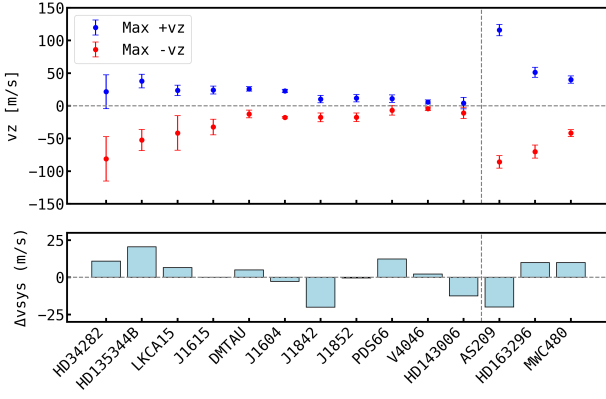


Figure 4. Top: maximum upward / downward ^{12}CO vertical velocities for the exoALMA targets to which we apply a 1D analysis. On the right are the three MAPS targets. Bottom: differences in v_{sys} computed from the ^{12}CO line and CS (for exoALMA) and from ^{12}CO and C^{18}O (for MAPS).

Table 1

Maximum Positive and Negative ^{12}CO v_z , v_{sys} Values Estimated from ^{12}CO and CS for exoALMA Targets (This Work) and from ^{12}CO and C^{18}O for MAPS (A. F. Izquierdo et al. 2023). Values Are in m s^{-1} .

1D Analysis—exoALMA				
Target	$\uparrow v_z$	$\downarrow v_z$	$v_{\text{sys}}^{12\text{CO}}$	CS
HD34282	21.5 ± 26.0	-81.3 ± 34.0	-2326.3	-2337.3
HD135344B	37.8 ± 10.4	-52.5 ± 16.2	7086.5	7066.0
LkCa15	23.6 ± 7.9	-41.6 ± 26.5	6288.0	6281.5
J1615	24.0 ± 6.0	-32.5 ± 11.9	4751.0	4751.0
DM Tau	25.8 ± 3.5	-12.5 ± 6.0	6033.6	6028.6
J1604	22.7 ± 2.6	-17.7 ± 1.3	4618.4	4621.0
J1842	10.2 ± 5.7	-17.7 ± 6.8	5940.6	5960.7
J1852	11.6 ± 5.5	-17.5 ± 6.6	5472.6	5473.1
HD143006	4.0 ± 8.8	-10.9 ± 8.7	7715.9	7728.4
PDS66	10.9 ± 5.8	-6.9 ± 7.5	3964.2	3951.7
V4046 Sgr	5.8 ± 3.2	-4.5 ± 2.6	2929.8	2927.6
1D Analysis—MAPS				
Target	$\uparrow v_z$	$\downarrow v_z$	$v_{\text{sys}}^{12\text{CO}}$	C^{18}O
AS209	115.9 ± 8.7	-86.0 ± 9.6	4640.0	4660.0
HD163296	51.2 ± 7.7	-70.2 ± 10.1	5770.0	5760.0
MWC480	40.1 ± 5.7	-41.8 ± 5.4	5100.0	5090.0
Spirals—exoALMA				
Target	$\uparrow v_z$	$\downarrow v_z$	$v_{\text{sys}}^{12\text{CO}}$	CS
MWC758	~ 500	~ -265	5894.4	5899.8
CQ Tau	~ 352	~ -290	6186.9	6180.2
SY Cha	~ 259		4103.6	4113.5

and 20 m s^{-1} (AS209; A. F. Izquierdo et al. 2023). These values provide some estimates of the potential systematic offsets affecting each v_z measurement discussed in this work. Applying such offsets in v_z , targets such as HD34282 and HD135344B exhibit even stronger downward motions, up to 92 ± 34 and $72 \pm 16 \text{ m s}^{-1}$, respectively, while J1842 and AS209 show enhanced upward motions, up to 30 ± 6 and $136 \pm 8 \text{ m s}^{-1}$, respectively. Targets like PDS66 and HD143006 would have a constant downward and upward flow, of about 10 m s^{-1} , respectively. However, for most targets, the offset in v_{sys} is rather marginal and does not affect our interpretation of the direction of vertical motion. We

acknowledge the limits of this simple analysis, as the measured offset traces the intrinsic difference in vertical motions between lines tracing different heights, which are very much dependent on the ongoing physical processes and the vertical thermal structures of the disks. However, we use the best currently available observational datasets to estimate the reliability of our derived vertical motions. Independent from the offset, the peak-to-trough values for each target represent the maximum amplitudes of vertical motion that can be measured at those targets, which in ^{12}CO range from $\sim 30 \text{ m s}^{-1}$ (J1852) to $\sim 100 \text{ m s}^{-1}$ (HD34282) in the exoALMA sample and from 80 m s^{-1} (MWC480) to 200 m s^{-1} (AS209) in the MAPS targets.

5.2. Radial Profiles

As explained in Section 3.2, we compute the radial profiles of the azimuthally averaged vertical velocities for 11 targets. Figures 5 and 6 show these profiles for the eight targets exhibiting significant upward or downward motion in the ^{12}CO line, defined as cases where the maximum velocity exceeds the standard deviation along the azimuth. The corresponding ^{13}CO profiles are also shown. For completeness, as the δv_ϕ profiles are reported in J. Stadler et al. (2025), we provide the v_r components for both lines in the lower panels of the same figures. In Appendix C, we present the v_z and v_r profiles for PDS66, V4046 Sgr, and HD143006—targets showing little to no vertical motion (Figure C2). As noted earlier, correcting for a potential offset due to an inaccurate v_{sys} would yield consistently low-velocity flows. The same profiles are provided for the MAPS targets (Figure C3). In all figures, the plotted error bars represent the azimuthal standard deviation at each radius, computed from the line-of-sight velocity maps and from the 2D v_r maps (see Figure C4 for the latter, whose complex structure interpretation is beyond the scope of this Letter). In Figures 5, 6, C2, and C3, we also show the ^{12}CO folded residual maps, to the right of the radial profiles. They provide a useful visual diagnostic, as coherent vertical motions appear as azimuthally extended arcs of a single color, as illustrated in Figure 1. The clearest examples of such coherent flows in the folded residual maps are J1604, J1852, and DM Tau.

For the following discussion, we consider the radial profiles as retrieved from our analysis, since the offset in v_{sys} does not affect our conclusions. Looking at the ^{12}CO radial profiles, we find that HD135344B, J1615, J1604, and DM Tau, as well as the three MAPS targets, show vertical motions with values above three times the azimuthal standard deviation. HD34282, LkCa15, J1842, and J1852 also show vertical motions, although with values above twice the azimuthal standard deviation. We note that V4046 Sgr shows a small upward flow at 160 au (5.8 m s^{-1}).

The vertical velocity profiles exhibit two types of behavior. The first type of v_z modulation is characterized by a single transition from downward to upward motion, with a broad downward plateau that reverses into an upward plateau at larger radii. We note that there is no case where the transition occurs from upward to downward. This behavior is observed in J1604 and J1852, with the transition occurring at approximately 179 and 150 au, respectively. AS209 also exhibits this pattern, with the transition point located at ~ 157 au.

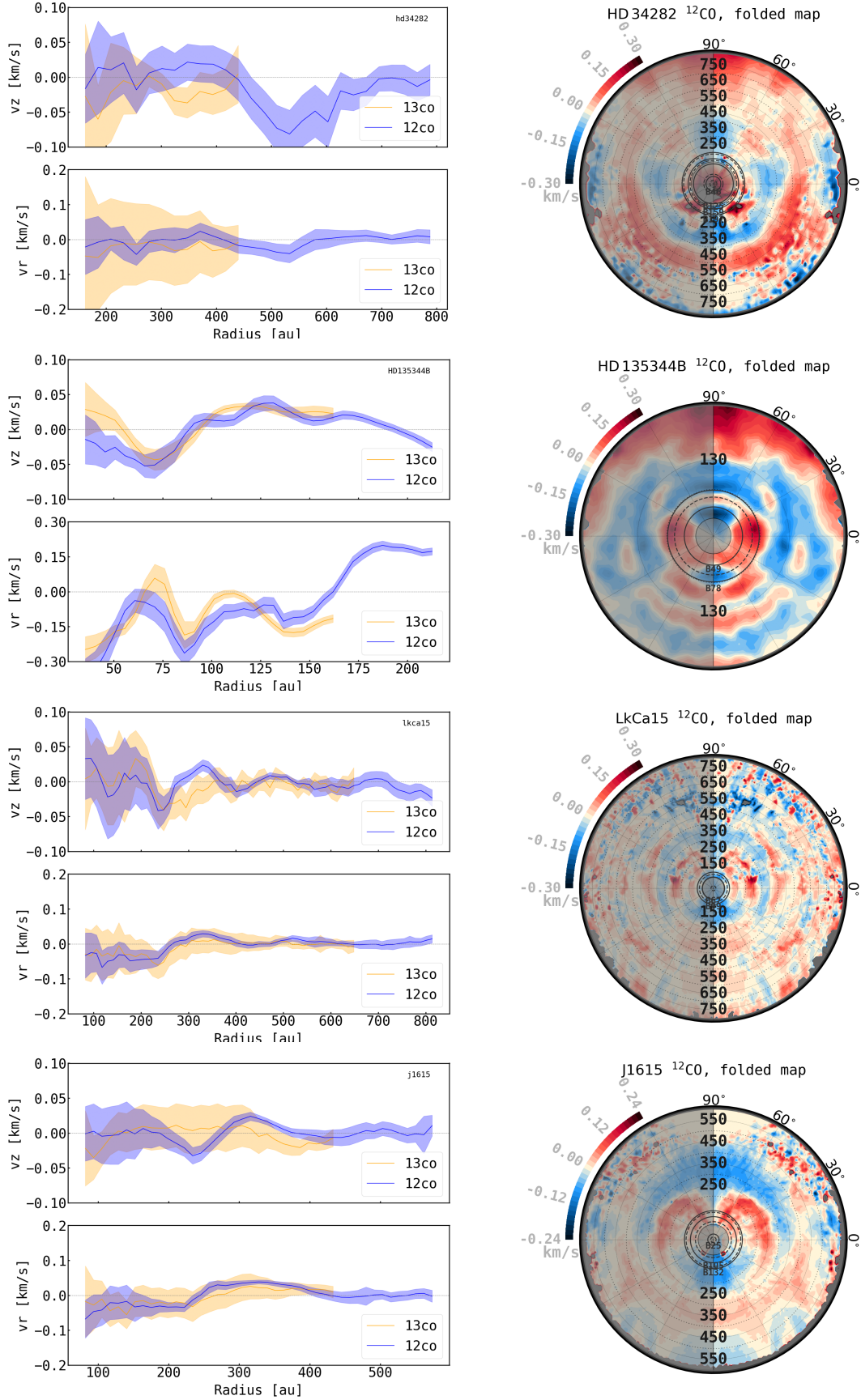


Figure 5. Gallery of azimuthally averaged vertical (top panels) and radial velocity (bottom panels) profiles. The blue lines are ^{12}CO and the yellow lines are ^{13}CO . The error is the standard deviation along each annulus. On the right, the folded residual maps for ^{12}CO are presented, in which any significant coherent vertical flow appears as an arc of uniform color (blue or red).

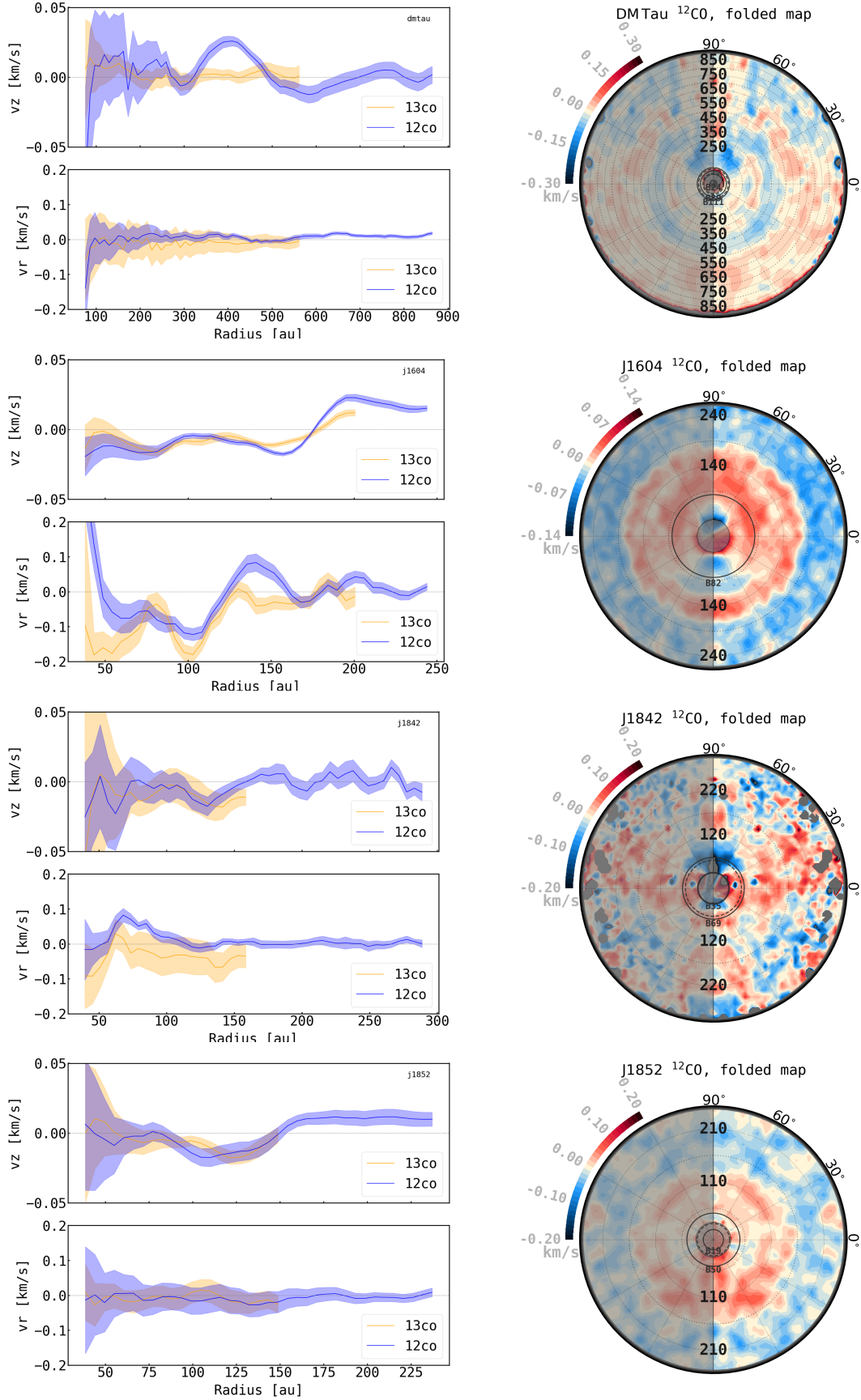


Figure 6. The same as Figure 5. The folded residual maps for ^{12}CO (on the right) show any significant coherent vertical flow appearing as an arc of uniform color (blue or red). DM Tau, J1604, and J1852 provide the clearest examples.

The second, observed in the majority of the disks, such as LkCa15, J1615, and DM Tau, is alternating upward and downward motions. To characterize this behavior, we identify peaks and valleys in the profiles using the `scipy.signal.find_peaks` function. We discard any peaks and valleys whose error bars make the feature consistent with zero. The radial locations of the peaks and valleys for the v_z profiles are shown in Table A2 and will be discussed further in Section 6.1. For instance, LkCa15 displays two upward- and two downward-velocity “bumps,” with zero-crossing points at approximately 280, 368, 445, and 522 au—that is, spaced nearly every ~ 80 au (i.e., about two beams). DM Tau exhibits upward motion from 322 to 494 au, with the peak at 410 au, followed by downward motion from 494 to 697 au, with the valley at 583 au. Similarly, J1615 shows downward motion from 187 to 275 au, followed by upward motion extending from 275 to 380 au. Among the MAPS sources, HD163296 and MWC480 display comparable “oscillatory” v_z behaviors, as reported previously (R. Teague et al. 2021; A. F. Izquierdo et al. 2023). In particular, MWC480 reveals multiple alternating upward and downward motions spanning the full radial extent of the disk, spaced by approximately 30, 40, 50, and 70 au. In contrast, HD34282 and J1842 only show a single broad downward-velocity bump, radially extending over 265 au and 48 au, respectively. Regardless of the specific radial profile behavior, the amplitudes of the detected vertical motions in these cases remain on the order of a few tens of m s^{-1} . All together, our results indicate that upward motions, possibly tracing winds, that span the entire radial range are not commonly observed in the ^{12}CO emission but instead tend to dominate in the outer regions of the disk.

The ^{13}CO radial profiles are generally noisier and extend over a smaller radial range, reflecting the lower signal-to-noise ratio of the data cubes. Nevertheless, in certain targets, including J1604, HD135344B, J1852, and LkCa15, the ^{13}CO profiles exhibit comparable amplitudes and similar radial structures to those seen in ^{12}CO , suggesting that the vertical layers traced by the two lines undergo similar vertical motions. For HD135344B, the vertical velocity crosses zero at approximately the same location for both ^{12}CO and ^{13}CO .

We also note that in some targets—such as HD34282, J1615, LkCa15 (Figure 5), and, most strikingly, MWC480 (Figure C3)—there is a positive correlation between v_z and v_r motions, indicating the presence of upward+outward or downward+inward flows. Interestingly, HD163296 shows an anticorrelation between v_z and v_r , suggesting instead a combination of upward+inward and downward+outward motions. These patterns may point to different circulation or transport mechanisms at play in these disks.

5.3. Spiral Targets

As discussed in Section 3.2, to quantify the vertical motions in targets exhibiting prominent spiral structures in the line-of-sight velocity residuals, we do not perform any azimuthal averaging of v_z but instead use FILFINDER to trace the velocities along the spiral spines. This approach ensures a more accurate characterization of the vertical velocities at a given radial separation and prevents the underestimation typically caused by azimuthal averaging.

In MWC758, we identify two prominent spiral structures in the velocity residuals, as illustrated in Figure 7. Moving outward from the inner disk, the first spiral (top panel) traces

blueshifted residuals corresponding to upward motions, extending from the inner regions out to approximately 165 au, with peak velocities reaching 135 m s^{-1} . The second spiral (second panel) traces redshifted residuals indicative of downward motions, spanning from roughly 60 au to 230 au, with velocities up to -265 m s^{-1} . The redshifted spiral shows clear azimuthal dependence, with higher velocities along the eastern side of the disk. To highlight the necessity of this approach over simple azimuthal averaging, Figure C5 compares the two methods for MWC758.

In addition to these spirals, we detect an outer annulus (third panel) characterized by strong blueshifted velocities, which we interpret as a signature of a disk wind. This component extends to the outermost detectable radius in our data and exhibits velocities up to 500 m s^{-1} , significantly higher than those observed in any other exoALMA or MAPS target. This filament seems to extend slightly farther, in a spiral-like shape, as seen in lower-resolution ^{12}CO cubes (see Figure 1 of M. Galloway-Sprietsma et al. 2025). We find that the velocity residual spirals exhibit significant width (as in, transverse to the filament). Both the blue- and redshifted spirals are also detected in the ^{13}CO velocity residuals, albeit with lower amplitudes, reaching approximately $\pm 100 \text{ m s}^{-1}$ (Figure 7, fourth and fifth panels).

Figure 8 shows the filaments extracted for CQ Tau and SY Cha. CQ Tau also shows two spirals, blue and redshifted as in MWC758. From the inside out, the first spiral (top left panel), characterized by blueshifted residuals, corresponds to upward motions extending from the inner disk to the outermost detectable radii, with peak velocities reaching 352 m s^{-1} in the innermost regions. The second spiral, traced along redshifted residuals, extends from 100 au to the outer disk, with peak downward velocities of -290 m s^{-1} . In addition to the spirals, we observe a Doppler flip in the southeast region. The velocity extraction for SY Cha reveals a single blueshifted filament, extending out to the radius where the diffuse emission begins (M. Galloway-Sprietsma et al. 2025), with peak velocities of 259 m s^{-1} . The velocity extraction for SY Cha should be interpreted with caution, due to the lack of a clearly identifiable redshifted spiral counterpart and the complex structure of the velocity residuals.

For these three targets with spiral features in their residual maps, we find differences in the systemic velocity between ^{12}CO and CS ranging from -9.9 m s^{-1} (SY Cha) to 6.7 m s^{-1} (CQ Tau), indicating only a minor offset relative to the amplitudes of the blue- and redshifted velocities measured.

6. Discussion

6.1. Correlation with Surface and Intensity Profiles

We compare the perturbations in the azimuthally averaged velocity profiles, v_z and v_r , to the molecular emission surface and temperature profile (traced through the peak intensity) structures identified in M. Galloway-Sprietsma et al. (2025). For this analysis, we focus on the disks for which we run an axisymmetric analysis (Figures 5, 6, and C2) and do not include disks with large-scale spirals. HD135344B and J1604 being nearly face-on, no surface could be retrieved, and we exclude these two disks from our analysis, even though they both show vertical motions. We also exclude PDS66, which shows a flat surface. This leaves a sample of nine disks. As mentioned previously, the radial locations of the peaks and

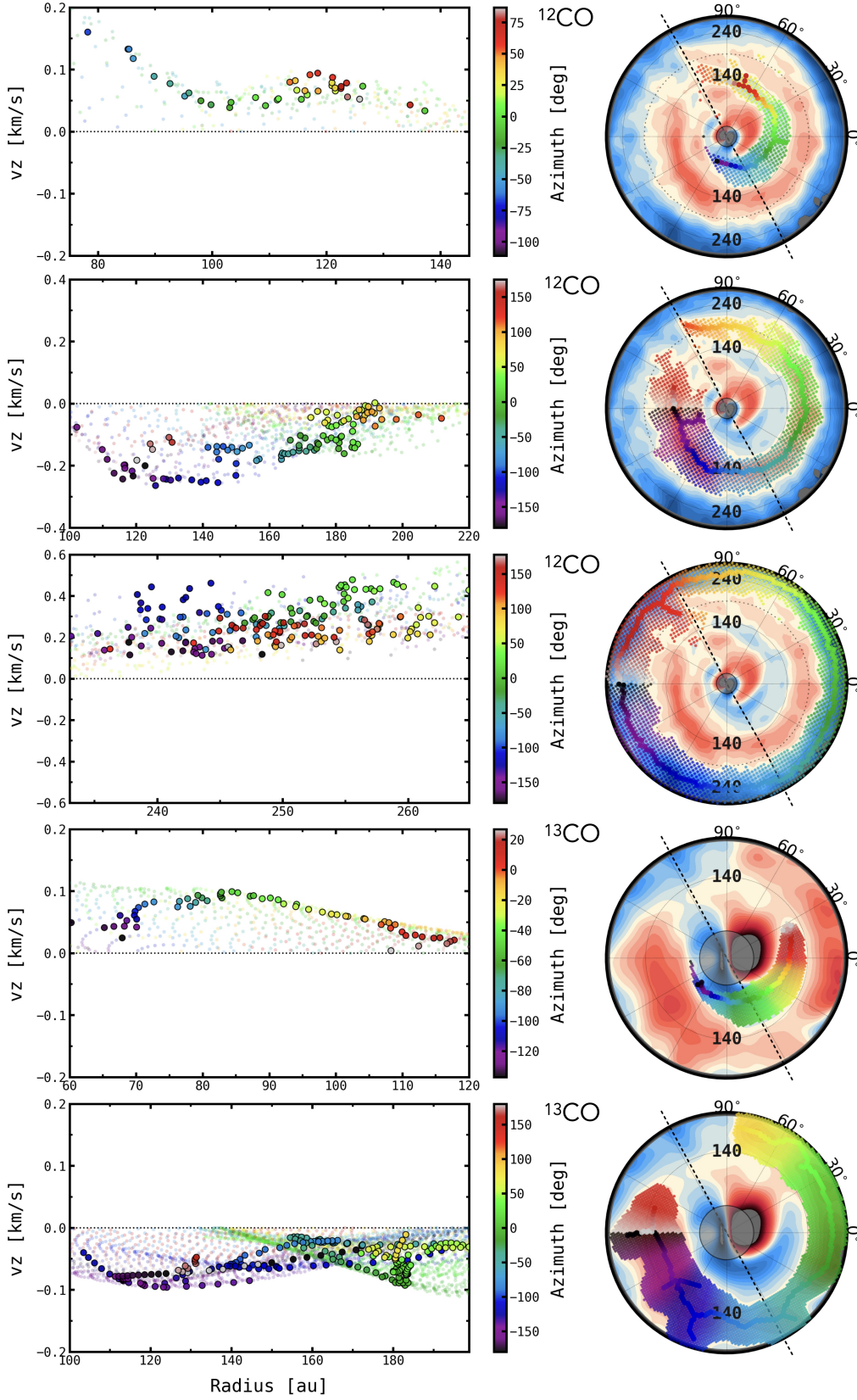


Figure 7. Line-of-sight velocities, which we interpret as vertical motion v_z , measured along filaments for MWC758 in ^{12}CO (top three panels) and ^{13}CO (bottom two panels). The dark colored points are the velocities extracted along the spine of the filaments depicted in the residual maps, presented to the right of each panel, while the pale dots trace the velocities measured along the width of the filaments. Note the different y-axis scales in the panels and the smaller radial extent of the residual maps in ^{13}CO .

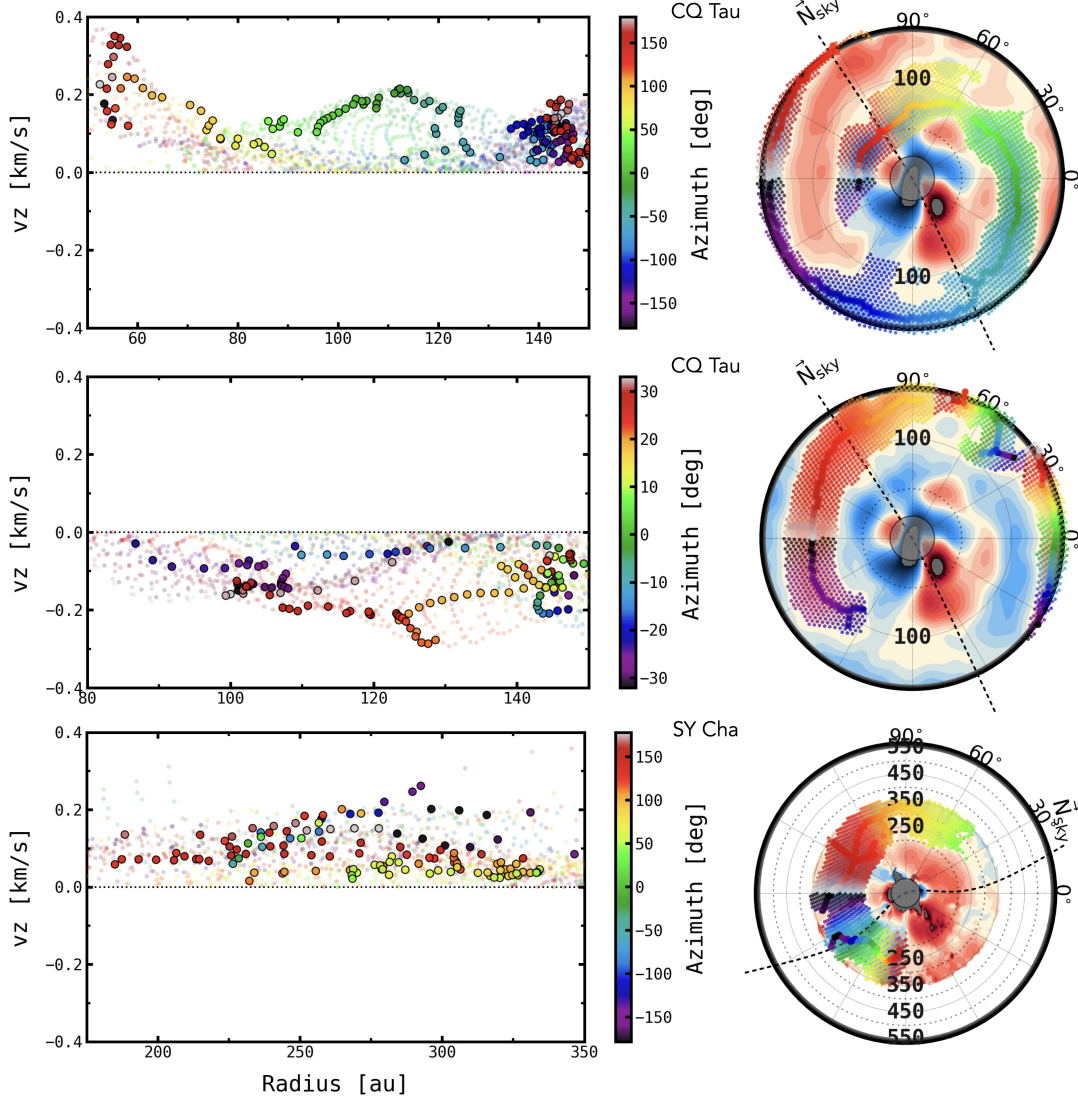


Figure 8. The same as Figure 7 but for CQ Tau and SY Cha in ^{12}CO . These targets do not show clear spiral counterparts in ^{13}CO .

valleys for the v_z profiles are shown in Table A2. We find that in over half of the sources (six out of the nine analyzed), perturbations in v_z spatially correlate within a beam size with substructures in either the emission surfaces or peak intensity profiles in one of the two lines. We illustrate this in Figure 9, by showing the ^{12}CO v_z as vertical arrows, overlaid on the nonparametric ^{12}CO surfaces by M. Galloway-Sprietsma et al. (2025). The peak intensity substructures found in ^{12}CO and ^{13}CO in the same work are also indicated. The arrows are given in units of the local sound speed, computed at the ^{12}CO emission layer by using the peak intensity values.

We note that the presence of substructures in the surface layer does not significantly bias the extraction of vertical motions measured on smooth parametric surfaces. Such surface perturbations are mirror-symmetric and disappear when folding residual maps and averaging, as in our methodology (see A. F. Izquierdo et al. 2023, their Figure B15; A. F. Izquierdo et al. 2026). In addition, M. Galloway-Sprietsma et al. (2023) found that there were no significant changes in the velocity residuals of AS209 when comparing a smooth surface to a structure surface.

The observed correlation is most pronounced in LkCa15, which exhibits multiple vertical velocity perturbations in ^{12}CO .

The upward motion at 330 au is coincident with an emission surface gap at 342 au and a peak intensity bump at 330 au. The upward motion at 472 au is also colocated with the peak intensity perturbations seen in both ^{12}CO and ^{13}CO . J1615 has a prominent vertical velocity structure at 234 au (downward motion), colocated with a ^{13}CO surface gap, and at 316 au (upward motion), colocated with a dip in the ^{12}CO peak intensity. In V4046 Sgr, the only vertical velocity peak at 161 au is spatially coincident with a prominent surface gap seen in both ^{12}CO and ^{13}CO .

Three disks have surface and/or intensity substructure correlated with radial velocity profile perturbations. In all three cases, the colocated structure is associated with the v_r profile crossing zero. In LkCa15, the dips in the peak intensity are at the same locations where v_r crosses zero (~ 270 au and ~ 450 au). The same is seen in J1842 at ~ 170 au and in J1615 at ~ 220 au.

Spatial correlations between velocity gradients and emission surface structures have been reported in previous studies,

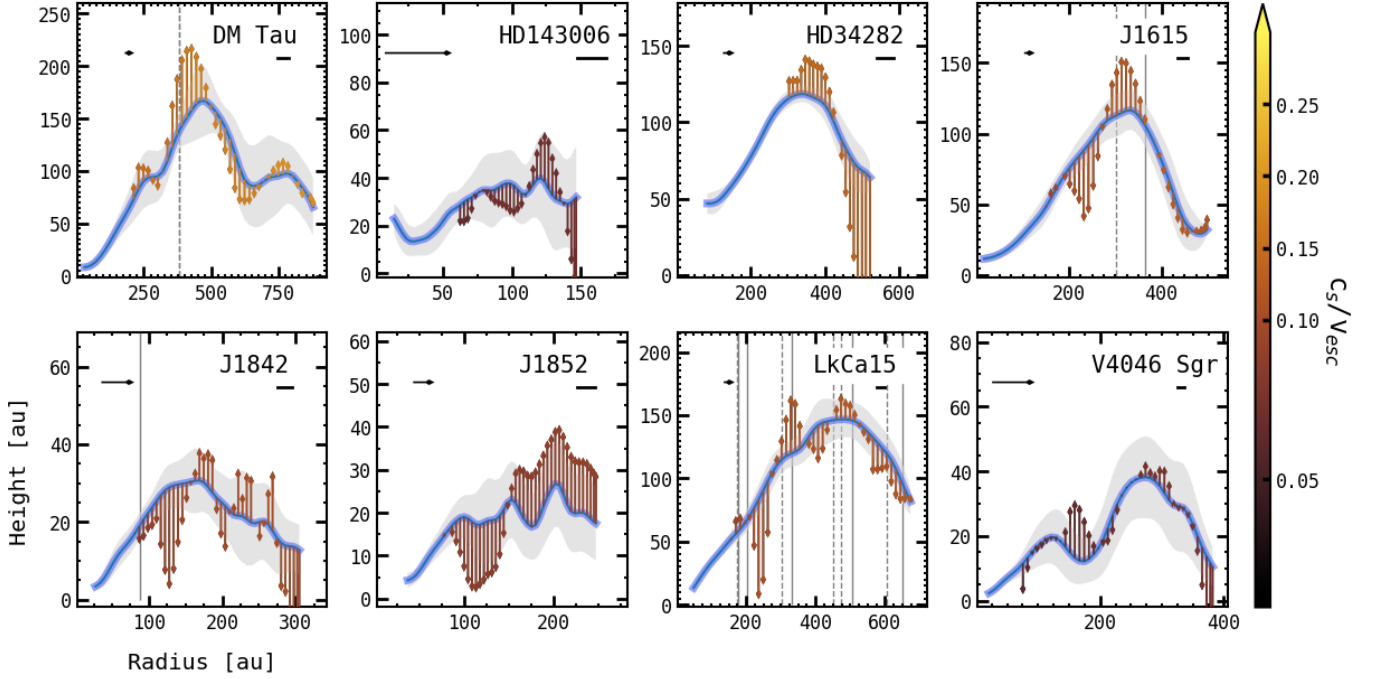


Figure 9. Vertical gas flows (arrows) along the ^{12}CO nonparametric surfaces (blue). The v_z arrows are scaled to the local sound speed at each radius, where the reference arrow in the top left corner represents $0.1c_s$. The arrows are colored by c_s/v_{esc} . The full (dashed) lines indicate the ^{12}CO and ^{13}CO intensity peaks (troughs) obtained from M. Galloway-Sprietsma et al. (2025). The beam sizes are shown in the top right corners.

notably in HD163296, MWC480, and AS209, which we reanalyzed in this Letter (R. Teague et al. 2019; A. F. Izquierdo et al. 2023; T. Paneque-Carreño et al. 2023). The variations in the emission surface likely reflect underlying changes in the local temperature structure and, potentially, surface density, as for example in the case of a gap. This is the case for AS209, for which M. Galloway-Sprietsma et al. (2023) speculated that winds driven by ambipolar diffusion are favored in the depleted regions where the gas density is very low. Focusing on the intensity and surface gaps observed in ^{12}CO , we find that DM Tau, J1615, and LkCa15 exhibit upward v_z associated with an intensity gap, similar to AS209, though at lower velocities and with less pronounced gaps. In contrast, J1615, J1842, and J1852 display downward v_z linked to a surface gap, while J1852, LkCa15, and V4046 Sgr show the opposite trend—that is, upward v_z associated with a surface gap.

6.2. Origin of the Vertical Motions

As illustrated in Section 4, planet–disk interactions and the three instabilities that we consider produce vertical motions of the order of $50\text{--}100\text{ m s}^{-1}$, with oscillatory behavior across the disk radial extent and stronger amplitude for the planet case at the edge of the gaps (e.g., J. Bi et al. 2021; X. Hu et al. 2025b). Winds, on the other hand, can produce high-amplitude vertical motions (e.g., X. Hu et al. 2025a). The amplitudes of the motions inferred from such models strongly depend on the assumed thermodynamics and vertical temperature structure (H.-G. Yun et al. 2025), as well as on the disk geometry itself, preventing the inferring of a single, coherent, mechanism as the origin of the vertical motions for each disk. In addition, the processes at play are not mutually exclusive, and multiple processes can occur at different locations of the disk.

For all targets, we compare the maximum vertical velocities to the local sound speed, c_s , and report the values in Table 2. Taking the absolute maximum values of v_z , we find that they range from a few percent to $\sim 25\%$ of the local sound speed (HD34282), reaching ~ 1.8 times the local sound speed for the outer wind in MWC758.

Fast vertical motions. MWC758, CQ Tau, and SY Cha show the strongest vertical motions, with non-Keplerian velocities appearing as spiral features and, in the case of MWC758, a fast wind in the outer disk. HD34282 and HD135344B exhibit the next-highest vertical velocities, both mostly downward motions. Interestingly, all four of these disks are around Herbig Ae/F stars, suggesting a possible correlation between the vertical gas motions and stellar mass. However, this trend is not consistent across the full exoALMA and MAPS sample: HD143006 (a G6 star) shows no significant vertical motion, while SY Cha or AS209, lower-mass K-type stars, display some of the most prominent vertical velocities.

It is also unclear whether the presence of strong vertical motions is directly related to the disk substructure, like cavities, gaps, and rings, as the measured motions extend beyond any continuum structure. Nevertheless, the four targets with strong vertical motions are transition disks with significant asymmetry in the continuum (P. Curone et al. 2025), spiral arms in infrared scattered light, and strong near-infrared excess (M. Benisty et al. 2023), indicating very dynamically active disks.

Spirals. While the current search for embedded companions within cavities has not yielded any candidate within the cavities of MWC758, CQ Tau, and SY Cha (e.g., B. B. Ren et al. 2023), the large spirals seen in the velocity residuals might be tracing indirect signatures of these companions. Particularly relevant for these two disks, E. Ragusa et al. (2024) studied the non-Keplerian features arising in a

Table 2
Maximum Upward and Downward v_z Values in ^{12}CO Compared to the Local Sound Speed c_s at the Corresponding Locations ($R_\uparrow$ and $R_\downarrow$)

Source	$\uparrow v_z/c_s$	$R_\uparrow$ (au)	$\downarrow v_z/c_s$	$R_\downarrow$ (au)	Pattern
HD34282	0.05 ± 0.06	348	0.25 ± 0.10	533	Single broad $\downarrow$ region
HD135344B	...	132	...	66	Oscillatory v_z
LkCa15	0.08 ± 0.03	330	0.14 ± 0.08	236	Oscillatory v_z
J1615	0.08 ± 0.02	316	0.10 ± 0.04	234	Oscillatory v_z
DM Tau	0.11 ± 0.02	410	0.06 ± 0.03	583	Oscillatory v_z
J1604	...	201	...	157	Single transition $\downarrow$ to $\uparrow$
J1842	0.03 ± 0.02	266	0.05 ± 0.02	130	Single broad $\downarrow$ region
J1852	0.04 ± 0.02	215	0.05 ± 0.02	110	Single transition $\downarrow$ to $\uparrow$
HD143006	0.01 ± 0.03	125	0.02 ± 0.03	157	Weak/no v_z
PDS66	...	103	...	44	Weak/no v_z
V4046 Sgr	0.02 ± 0.01	161	0.02 ± 0.02	370	Weak/no v_z
MWC758	~ 1.82	>220	...		$\uparrow v_z$ [wind]
CQ Tau	~ 0.70		~ 0.63		Spirals $\uparrow$ and $\downarrow$
MWC758	~ 0.26		~ 0.73		Spirals $\uparrow$ and $\downarrow$
SY Cha	~ 0.70		...		Spirals $\uparrow$ and $\downarrow$

Note. For disks for which a surface could not be extracted in M. Galloway-Sprietsma et al. (2025), we do not derive any v_z/c_s . For the spirals, we do not report any error, as we extract the maximum velocity along the spiral spines, while for the azimuthal averaged v_z , the error on v_z represents the standard deviation along the azimuth and is propagated here, along with the standard deviation of the brightness temperature from M. Galloway-Sprietsma et al. (2025). We also report the observed v_z pattern.

circumbinary disk with eccentric motions. They consider a massive companion with a mass of 10% of the central star and a semimajor axis of 15 au. Such a companion carves a 40 au wide cavity, with an eccentricity of 0.2, which in turn induces disk eccentricity, as commonly seen in binary–disk interactions (e.g., Z. Regály et al. 2011; R. Miranda et al. 2017; E. Ragusa et al. 2020). The change in the vertical component of gravity along the eccentric orbit produces vertical oscillations in the disk, inducing changes in the disk’s local scale height. This oscillatory motion gives rise to strong vertical velocities, up to 400 m s^{-1} , and leads to two spiral arms, blue- and redshifted, strikingly similar to what is observed in MWC758 and CQ Tau. They also found that the amplitude of the vertical motions scales with height, which supports a stronger v_z for ^{12}CO than for ^{13}CO . While the range of disk and companion parameters that would lead to such eccentric and vertical motions is too large to provide narrow constraints on the nature of the companions possibly embedded within the cavities of MWC758 and CQ Tau, this work provides strong support for their presence. An alternative explanation for the observed spirals is that they are tracing the *inner* spirals driven by an outer companion. This idea was originally proposed by R. Dong et al. (2015) to reproduce the infrared-scattered-light spirals in MWC758, as a planetary mass perturber in the cavity would induce *outer* spirals, which would be too tightly wound to match the pitch angle of infrared spirals. However, the vertical motions in the inner spirals induced by massive planets do not show strong enough amplitudes to match our observations (I. Rabago & Z. Zhu 2021), compared to the scenario of an eccentric stellar companion (E. Ragusa et al. 2024). Further support for an eccentric companion in the cavity comes from the measured eccentricity of the dust cavity in MWC758 ($e \sim 0.1$; R. Dong et al. 2018) and the tentative detection of spiral arms in the continuum of CQ Tau (F. Zagarria et al. 2025). Interestingly, another object in our sample, HD135344B, shows similar disk properties but does not show such high-velocity vertical motions. The differences might

originate in the nature (planetary versus stellar) and orbital parameters of the companion.

Large non-Keplerian-velocity $m = 2$ spirals can also be induced by temperature variations in the disks, induced by shadows. S. Zhang & Z. Zhu (2024) and A. Ziampras et al. (2025) investigated the dynamical impacts of shadows in transition disks that have misaligned inner disks and found that as the shadowed regions are cooler, asymmetric pressure gradients are generated, leading to spiral density waves detectable in their kinematical signatures with velocities about 100 m s^{-1} . However, the MWC758 and CQ Tau residuals show $m = 1$ spirals, hence this scenario does not appear to be likely.

Large-scale infall can also trigger spiral arms with large opening angles (P. Hennebelle et al. 2017; J. Calcino et al. 2025; C. Longarini et al. 2025) and result in disk eccentricity, naturally inducing vertical motions. While so far there are no indications of large-scale streamers in the vicinities of CQ Tau and MWC758, we note that in MWC758, the outer disk seems to extend slightly farther in a spiral-like shape in lower-resolution ^{12}CO cubes (see Figure 1 of M. Galloway-Sprietsma et al. 2025).

Radially varying v_z . For the targets analyzed with an axisymmetric method, as previously discussed, we see different behaviors. First, in both the exoALMA and MAPS targets (e.g., LkCa15 and MWC480), we detect oscillatory up-and-down motions with amplitudes of a few tens of m s^{-1} . This behavior is reminiscent of VSI simulations (see Figure 3; M. Barraza-Alfaro et al. 2025), although we cannot rule out multiple mechanisms coherently contributing to the observed vertical motions. However, the relatively low vertical velocities, below 10% of the sound speed at the emission heights traced by the ^{12}CO and ^{13}CO lines (roughly two to three scale heights), suggest that the disks are more dynamically quiet than expected from simulations. For example, large-scale upward flows associated with fully developed MRI can reach $v_z > 0.1 c_s$ at these heights (M. Flock et al. 2011). It is however very likely that the limited angular

resolution of our observations contributes to the low measured v_z , as small-scale turbulent eddies are washed out. A detailed comparison between the observed vertical motions and simulations is beyond the scope of this Letter. The second behavior, observed in, e.g., J1604, J1852, and AS209, is a transition from downward to upward vertical motions around 150 au, which would be consistent with the launching of a disk wind in the outer regions. J1615, J1842, and J1852 have downward motion associated with a surface gap, but J1604, HD34282, and J1615 show downward motion without clear counterparts in surface or intensity profiles. The origin of such downward motion over extended regions (100–200 au) is unclear.

Winds. We consider whether the upward motion may be driven by thermal pressure gradients. Fully ionized thermal winds would not be expected to have an appreciable molecular content (e.g., D. Hollenbach et al. 1993, 1994; R. D. Alexander et al. 2006). However, this does not need to be the case for winds launched from the outer disk. In this case, winds may be driven by X-ray heating, and advection can allow molecules to survive in the unbound gas (J. E. Owen et al. 2010; L. Wang & J. Goodman 2017; A. D. Sellek et al. 2024). Thermal winds are approximately launched when

$$c_s \gtrsim \sqrt{\frac{\gamma-1}{4\gamma}} v_{\text{esc}} \approx 0.3 v_{\text{esc}}, \quad (5)$$

where c_s is the sound speed, $v_{\text{esc}} = \sqrt{2GM_*/r}$ is the escape velocity, and assuming $\gamma = 5/3$ (K. Liffman 2003). We can then consider how close to meeting this criterion the gas at the CO emission surface is for the disks in our sample. In Figure 9, we show that in several of the disks, c_s/v_{esc} is substantially lower than 0.3. This would be in line with simulations showing that for Class II disks, little CO is expected in the molecular component of the wind, unlike H_2 , which is expected to be abundant (A. D. Sellek et al. 2024—supported by the evidence of H_2 winds in JWST data; K. R. Schwarz et al. 2025a, 2025b). Indeed, any PE or MHD-driven wind might be active at higher altitudes above four to five scale heights (R. Nakatani et al. 2018; O. Gressel et al. 2020) with much larger vertical velocities, although without contributing to the observed ^{12}CO emission. Even in the outer regions of MWC758, where vertical velocities reach 1.8 times the sound speed, c_s/v_{esc} is ≤ 0.1 , meaning that the material is likely located at the very base of a wind. However, in a couple of cases, most evidently DM Tau and HD 34282, the highest point of the emission surface exhibits upward flow with c_s/v_{esc} close to 0.3. The coincidence of the maximum emission height and strongest vertical velocity could plausibly exhibit CO at or close to the wind-launching location. No work has yet explored the vertical velocity of molecular CO in a steady-state PE wind—in particular, whether $\sim 100 \text{ m s}^{-1}$ velocities may be reached under certain conditions. Although A. D. Sellek et al. (2024) suggested minimal CO advection into the wind, a systematic exploration of the conditions and degree to which CO can survive in the wind region is motivated by our findings.

7. Summary and Conclusions

In this Letter, we present a comprehensive and uniform analysis of the kinematics of 14 exoALMA targets using their

^{12}CO and ^{13}CO emission lines, with the aim of characterizing their vertical motions, also including three MAPS targets for comparison. We find that the majority of the disks exhibit ubiquitous vertical motions, although their amplitudes and morphologies vary from one disk to another. We summarize our findings as follows:

1. We derive vertical velocities from the line-of-sight velocity residuals, after subtracting a smooth Keplerian disk model using DISCMINER. We validate our v_z retrieval technique using hydrodynamical models of planet–disk interactions and three classes of disk instabilities (VSI, GI, and MRI) and find that, in those simulations, the complex kinematics does not lead to artificial vertical motions. We therefore demonstrate that in these simulations, vertical velocities can be reliably retrieved down to amplitudes of about 10 m s^{-1} , even if the disk dynamical state is high.
2. Depending on the level of asymmetry observed in the line-of-sight residual maps of our disks, we adopt two different approaches. For targets displaying clear spiral residual patterns (MWC758, CQ Tau, and SY Cha), we extract velocities from the 2D residual maps directly along the spiral spines. In MWC758 and CQ Tau, we detect strong upward and downward motions, reaching up to approximately 300 m s^{-1} , which appear as red- and blueshifted spiral structures. We interpret these features as signatures of disk eccentricity induced by massive companions.
3. For the remaining targets, we follow the methodology employed in previous studies and compute azimuthally averaged radial profiles. In those disks, two general behaviors emerge: (1) a radially oscillatory pattern of upward and downward motions, which likely trace meridional flows induced by disk instabilities; and (2) a smooth radial transition from downward to upward motions, which we interpret as evidence of disk winds in the outer regions. In general, these disks exhibit maximum vertical velocity amplitudes ranging from 30 to 100 m s^{-1} , corresponding to approximately 1% to 25% of the local sound speed, except for MWC758, which has a high-velocity disk wind in the outer disk, with velocities up to 500 m s^{-1} —that is ~ 1.8 times the local sound speed. The ^{13}CO observations are noisier and cover a smaller radial extent than the ^{12}CO line data, but in several targets they display similar amplitudes and radial trends to those derived from the ^{12}CO line.

In conclusion, based on the analysis of 14 exoALMA targets and three MAPS disks, we find that vertical motions are ubiquitous, as revealed by the ^{12}CO and ^{13}CO emission lines, indicating meridional circulations in disks, although our angular resolution might not be high enough to catch small eddies. In objects that show strong dynamical sculpting, such as transition disks like MWC758 and CQ Tau, these vertical motions may provide an additional constraint for identifying the origins of their cavities. Strong molecular winds appear to be relatively rare in ^{12}CO and ^{13}CO , and the measured vertical motions might be probing the bases of such winds. Overall, the velocity structure traced in our disks is complex and highly diverse, suggesting that it is unlikely that a single physical mechanism is responsible for the vertical motions observed across all systems. This diversity motivates the need for further

theoretical and numerical studies capable of explaining the full range of observed kinematics.

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Appendix A Data Table

Table A1 gives the information on the ^{12}CO , ^{13}CO cubes used in this Letter, the method to compute the centroid map, the outer disk radius considered in the analysis, as well as the means and standard deviations of the corresponding folded maps. Table A2 gives the locations of the v_z peaks and dips for ^{12}CO and ^{13}CO .

Table A1
Data Cubes Used in This Letter

Target	^{12}CO Res.	Centroid	Rout Cut (% Rout)	Mean (km s^{-1})	Std (km s^{-1})	^{13}CO Res.	Centroid	Rout Cut (% Rout)	Mean (km s^{-1})	Std (km s^{-1})	Comment
MWC758	$0''.15$	G	0.80	0.25	0.18	$0''.30$	G	0.65	0.08	0.03	Spirals
SY Cha	$0''.30$	DB	1.0	0.22	0.18	$0''.30$	DB	1.0	0.32	0.13	Spirals
$R < 360$ au			0.6	0.20	0.14			0.80	0.315	0.13	
CQ Tau	$0''.15$	G	1.0	0.22	0.12	$0''.15$	G	0.85	0.18	0.10	Spirals
AA Tau	$0''.30$	DB	0.82	0.21	0.11	$0''.30$	DB	0.30	0.29	0.16	Excluded
HD135344B	$0''.15$	G	1.0	0.14	0.05	$0''.15$	G	0.95	0.12	0.03	1D
LkCa15	$0''.30$	DB	1.0	0.14	0.09	$0''.30$	DB	0.8	0.24	0.15	1D
HD34282	$0''.30$	DB	0.9	0.14	0.01	$0''.30$	DB	1.0	0.33	0.24	1D
J1842	$0''.15$	DB	0.92	0.12	0.07	$0''.15$	DB	0.50	0.11	0.05	1D
DM Tau	$0''.30$	DB	1.0	0.11	0.07	$0''.30$	DB	0.68	0.035	0.02	1D
J1615	$0''.30$	DB	0.95	0.10	0.07	$0''.30$	DB	0.85	0.14	0.10	1D
HD143006	$0''.15$	G	1.0	0.06	0.03	$0''.15$	G	0.92	0.07	0.02	1D
J1852	$0''.15$	G	0.9	0.05	0.02	$0''.15$	G	0.6	0.06	0.03	1D
J1604	$0''.15$	B	1.0	0.05	0.02	$0''.15$	B	1.0	0.05	0.02	1D
PDS66	$0''.15$	G	0.92	0.04	0.01	$0''.15$	G	0.55	0.13	0.06	1D
V4046 Sgr	$0''.30$	G	1.0	0.02	0.01	$0''.30$	G	0.95	0.03	0.01	1D
HD163296	$0''.15$	DB	0.95	0.21	0.13	...	...	...	...	...	1D
AS209	$0''.15$	G	0.82	0.25	0.10	...	...	...	...	...	1D
MWC480	$0''.15$	G	1.0	0.09	0.04	...	...	...	...	...	1D

Note. For each line, we provide the angular resolution and the method used to retrieve the centroid velocity (A. F. Izquierdo et al. 2025—G: Gaussian; B: Bell; and DB: double Bell). We also report the outer radii considered in our analysis in the fraction of the cube outer radius and the means and standard deviations computed in the corresponding folded residual maps. The last three rows correspond to the MAPS cubes. For SY Cha, we report two lines—the second one corresponding to the folded map after we cut off the diffuse emission in the outer disk.

Table A2
Radial Locations (in Astronomical Units) of the v_z Peaks and Dips for ^{12}CO and ^{13}CO

Source	^{12}CO		^{13}CO	
	Peaks	Dips	Peaks	Dips
DM Tau	410	583	...	...
HD135344B	132	66	122	71
HD143006	...	...	...	...
HD34282	...	533	...	185, 347
J1604	201	157	...	81
J1615	316	234	250	94, 386
J1842	266	130, 204	...	79, 142
J1852	...	110	...	121
LkCa15	330, 472	236, 401, 566	...	248
V4046 Sgr	161	...	...	...

Note. The bold numbers represent the peaks and valleys associated with a surface or intensity substructure, within a beam size. We exclude HD135344B, J1604, and PDS66 (see the text). We refer to Table 4 of M. Galloway-Sprietsma et al. (2025) for a complete list of surface and intensity structures.

Appendix B Beam Convolution Effect

In Figure B1, we illustrate the effects of radiative transfer and beam convolution on the simulations of the VSI presented in Section 4. We show, on the one hand, the vertical velocities extracted directly from the model at a given height, $Z/R = 0.24$ (black line), and, on the other hand, the velocities extracted from the postprocessed simulated image after radiative transfer, both at infinite angular resolution (the yellow line) and at the exoALMA angular resolution (the pink line).

This comparison highlights two key aspects. First, the model velocities at $Z/R = 0.24$ have higher amplitudes than those obtained after radiative transfer, even at infinite angular

resolution. This indicates that, in the radiative transfer postprocessed image, multiple contributions from velocities at different vertical layers around the $\tau = 1$ surface can smear out the observed signal.

Second, by comparing the infinite-angular-resolution case with the exoALMA resolution, we find that beam convolution at the current exoALMA resolution can significantly reduce the amplitudes of the intrinsic velocities. This effect would be enhanced for processes where multiple velocity components coexist at the same location, but it would be minimized when the flow has a single coherent direction, as in the case of wind motions. Overall, this suggests that the vertical motions measured in our data are likely lower limits on the true velocity amplitudes.

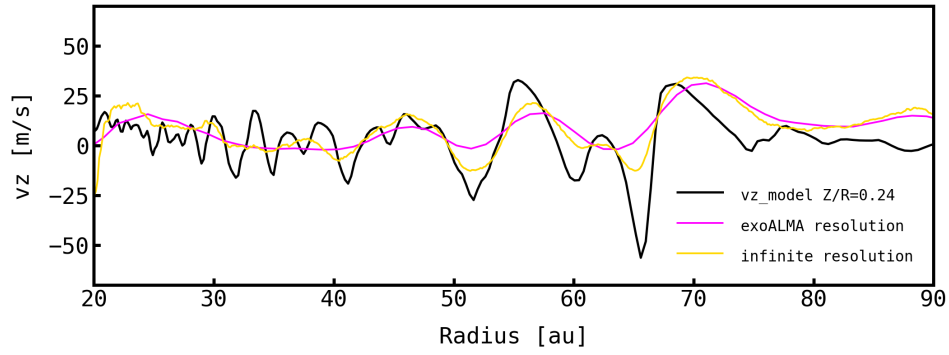
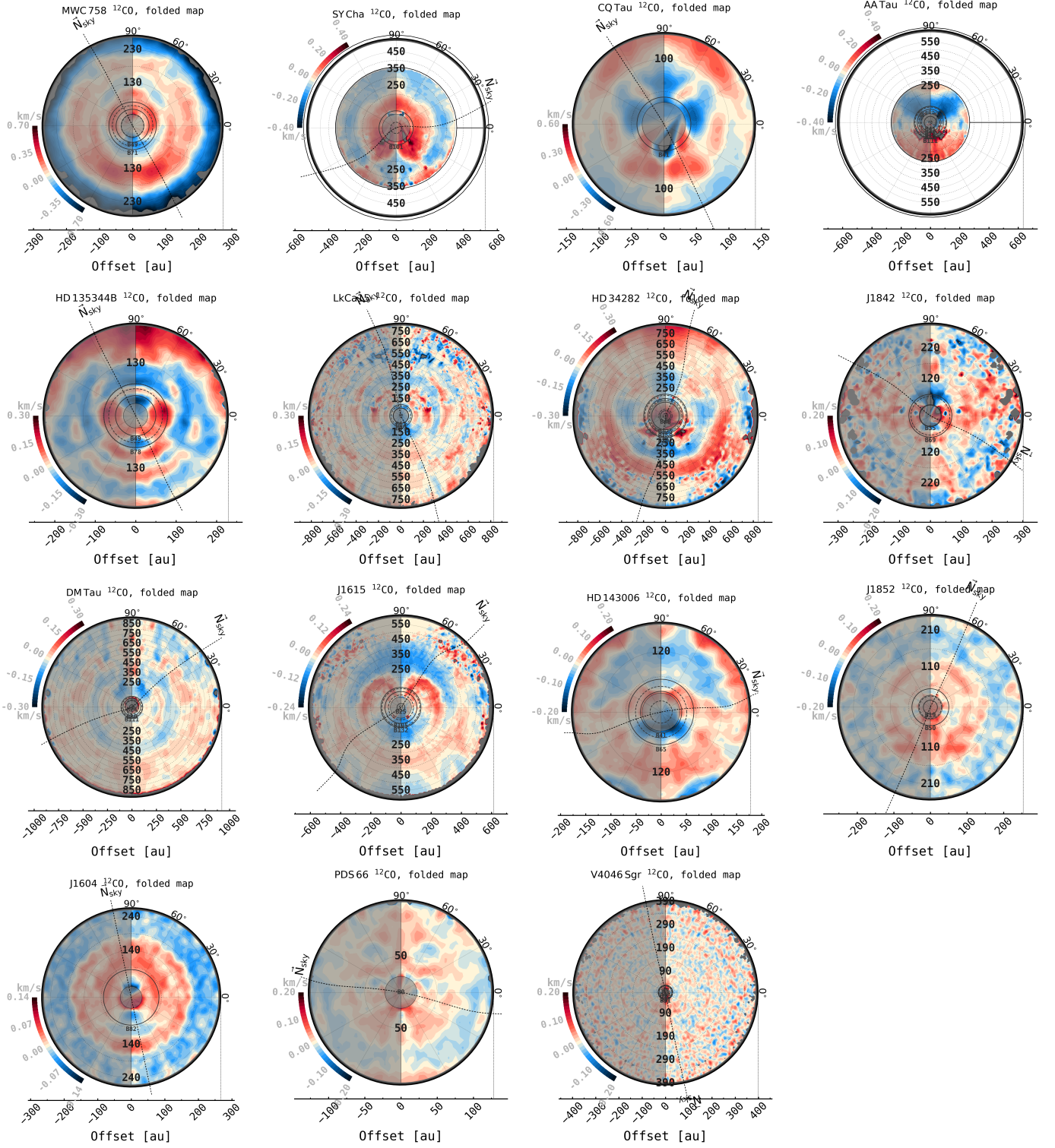


Figure B1. Comparison between the vertical velocities extracted from the VSI simulation at a given height with those extracted after postprocessing and beam convolution.

Appendix C Additional Figures

Figure C1 shows a gallery of folded residual maps. Figures C2 and C3 provide the radial profiles of vertical and radial velocities, for the exoALMA targets for which we do not detect clear

vertical motions and for the MAPS targets. Figure C4 shows a gallery of residuals in radial velocity components, from which an azimuthal profile is extracted. Figure C5 shows how a v_z extraction following an azimuthal average would underestimate the velocities in the spirals of MWC758.

Figure C1. Gallery of folded centroid residual maps of the ^{12}CO emission.

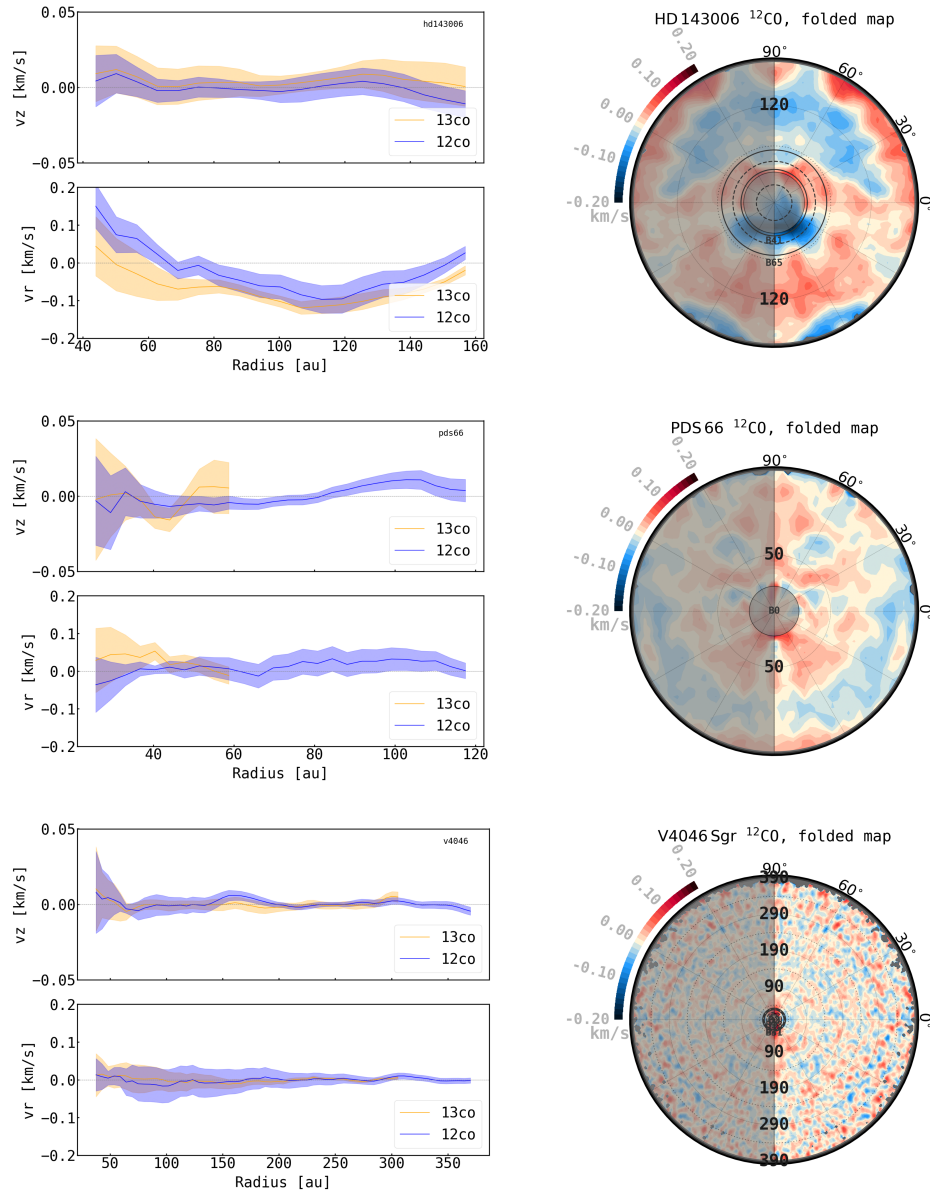


Figure C2. The same as Figure 5 but for three targets that do not show strong vertical motion.

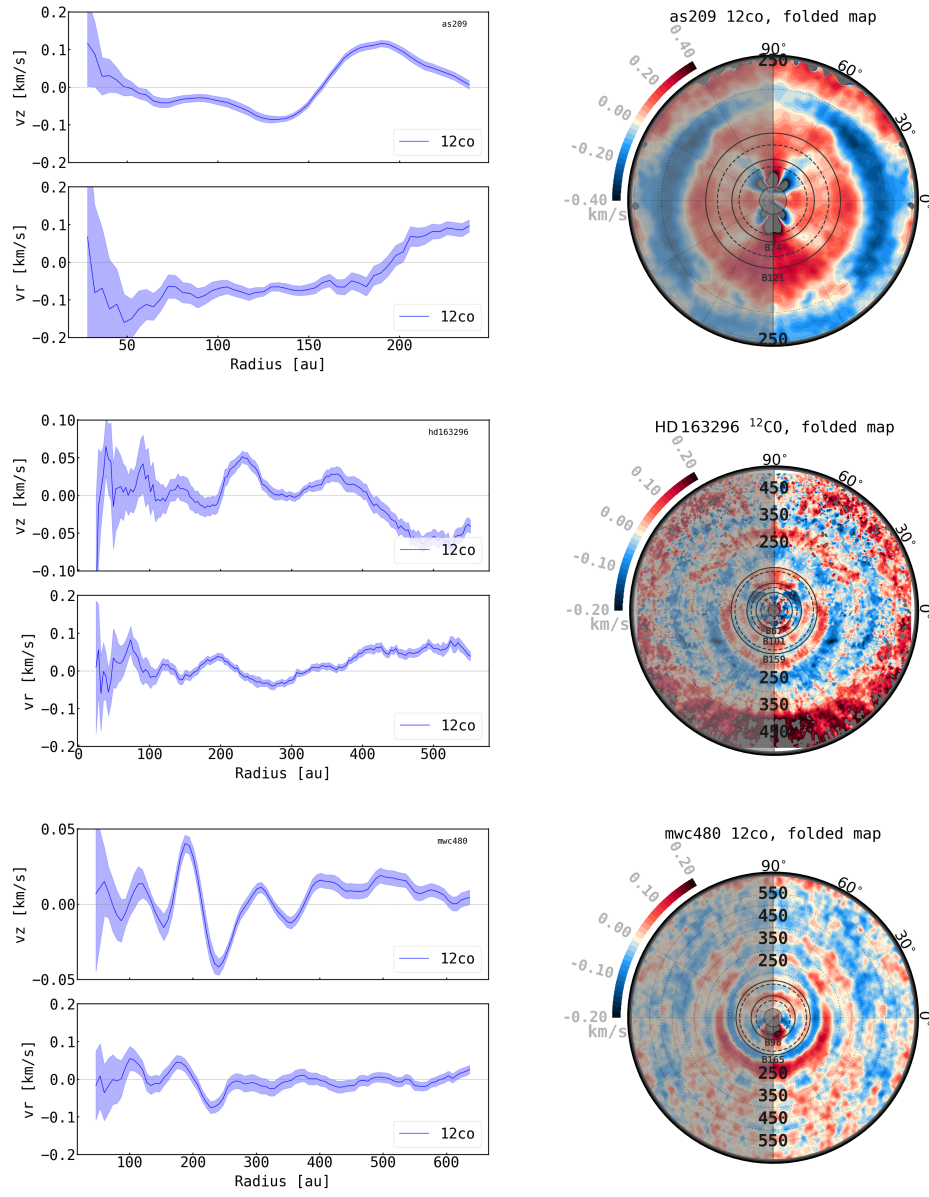


Figure C3. The same as Figure 5 but for MAPS targets.

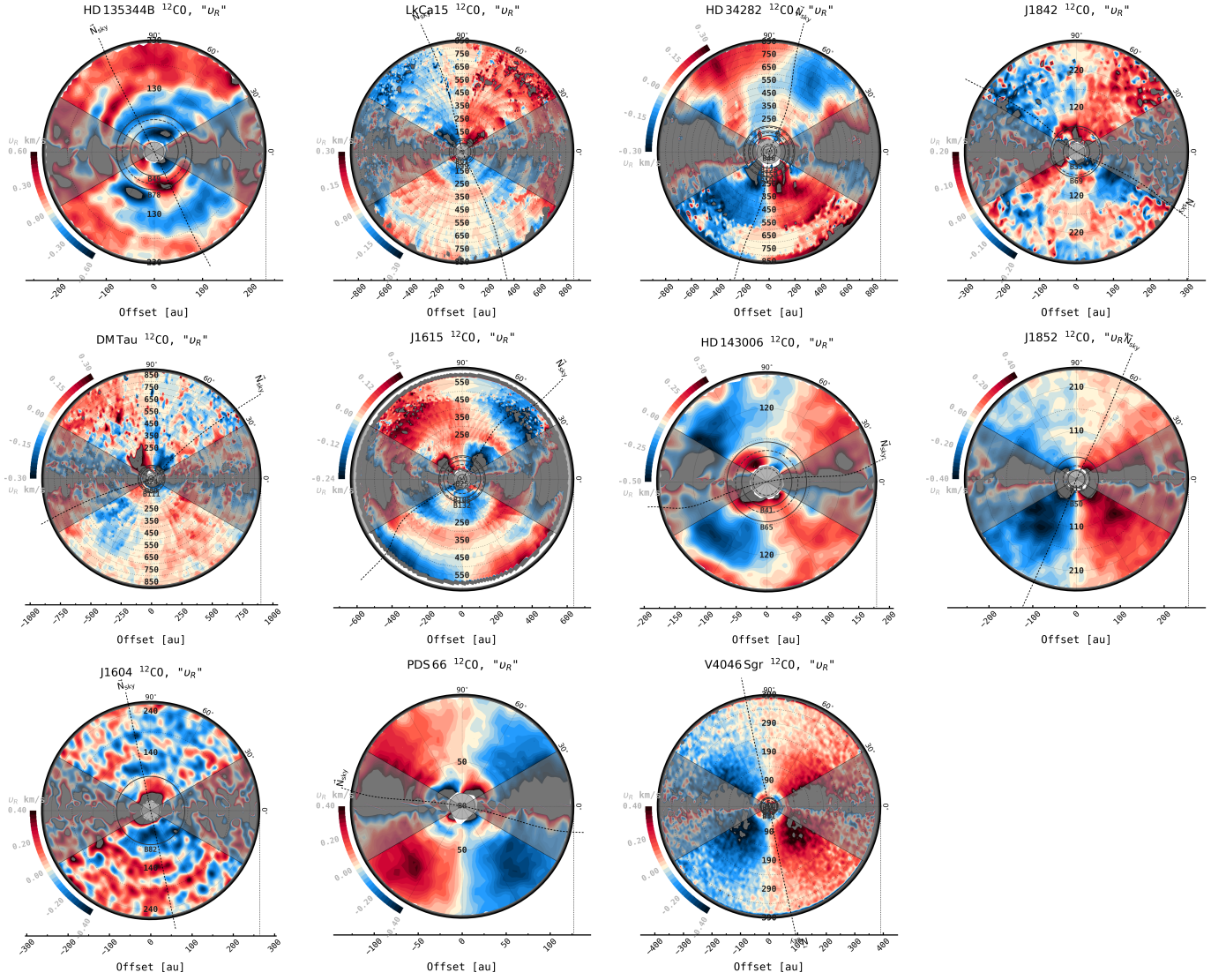


Figure C4. Gallery of 2D radial velocity residual maps for all targets for which we applied the 1D analysis obtained following A. F. Izquierdo et al. (2023). The gray area represents the major axis where the radial velocities diverge.

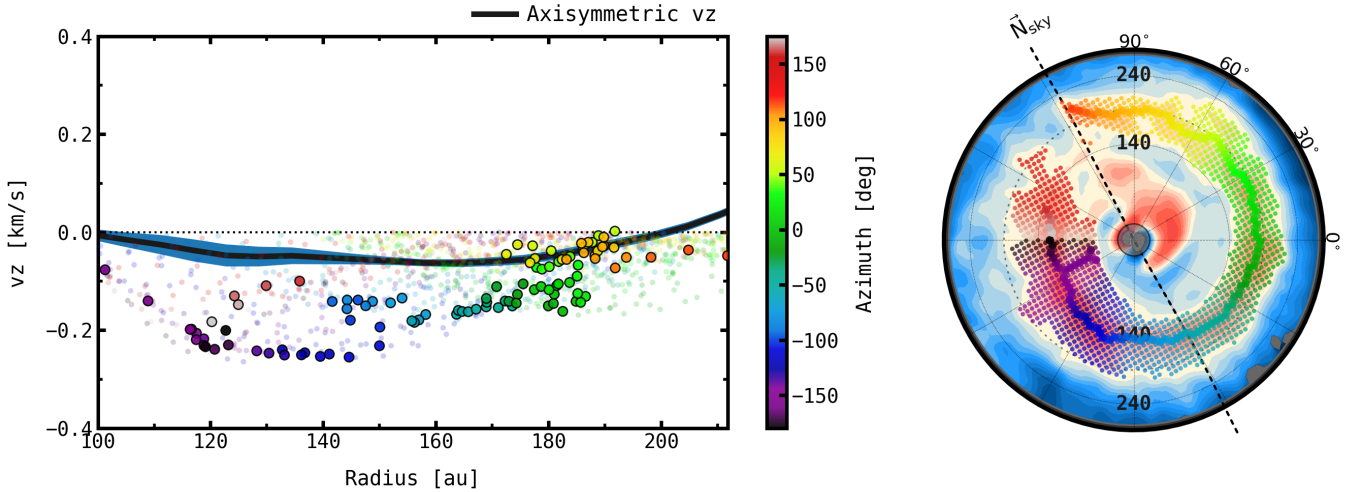


Figure C5. An axisymmetric analysis of the velocity in MWC758 (black line) would underestimate the vertical motions at the radial locations traced by the spiral spine.

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