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Isospin Symmetry Breaking Effects on the Mass-Radius Relation of a Neutron Star

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Abstract: Isospin symmetry breaking effects on the mass-radius relation of a cold, non-accreting neutron star are studied on the basis of two Skyrme Energy Density Functionals (EDFs). One functional contains isospin symmetry breaking terms other than those typically included in Skyrme EDFs while its counterpart is of standard form. Both functionals are based on the same fitting protocol except for the observables and pseudo-observables sensitive to the isospin symmetry breaking channel. The quality of those functionals is similar in the description of terrestrial observables but choosing either of them has a non-negligible effect on the mass-radius relation and tidal deformability of a neutron star. Further investigations are needed to clarify the effects of isospin symmetry breaking on these and other observables of neutron stars that are, and will become, available.

Keywords: isospin symmetry breaking; neutron star mass-radius relation; tidal deformability



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1. Introduction

The nuclear Equation of State (EoS) is one of the fundamental ingredients to describe the mass-radius relation in neutron stars [1–5]. The neutron star core is believed to cover the largest portion of the neutron star radius and mass. Despite the effect of the crust is non-negligible to properly determine such quantities (for more details see, e.g., Refs. [6,7]), relative effects might be safely studied neglecting the crust. It is customary to assume the core of a neutron star as made of β -stable neutrons, protons and electrons [8–13]. This allows easily calculating the mass-radius relation and compare to observational data in order to detect deviations from such an approximation. Different effects that may influence such a relation have been studied in the literature, e.g., three-neutron forces at high densities or the plausible appearance of new degrees of freedom such as hyperons or quarks. While models with hyperonic degrees of freedom predict a soft equation of state for the star, not being able to explain the largest masses observed to date (see, for example, [14]), models that account for a phase-transition to quark deconfined matter do not suffer from such problem [15].

Nuclear energy density functionals (EDFs) were proven to predict with a good accuracy the ground state as well as some excited state properties in terrestrial nuclei and have been also applied to predict the limits of nuclear existence [16,17]. The use of these models to the study of neutron-star matter may be regarded as strong extrapolation due to the difference in average density within the interior of a neutron star (over $2\rho_0$) and the interior of the atomic nucleus (ρ_0). Currently, EDFs are the only type of models that can be used to consistently predict the structure and composition of a neutron star outer core and of the neutron star crust at the same time [18]. Hence, to study how different physical effects may influence the mass-radius relation of a neutron-star by using EDFs is well justified.

Isospin symmetry breaking (ISB) effects, other than those originated from direct and exchange part of the Coulomb interaction, are not commonly included in the fitting of current EDFs. This is because assuming isospin symmetry of the nuclear strong interaction and neglecting Coulomb corrections beyond the direct and exchange terms is known to be a

very good approximation for the study of many observables in atomic nuclei. The most paradigmatic counter example is the excitation energy of the Isobaric Analog State (IAS) that has been measured for many nuclei with exquisite accuracy (cf. the compilation in Ref. [19]). The IAS energy depends only on ISB terms in the nuclear Hamiltonian. Specifically,

$$E_{\text{IAS}} = \frac{\langle 0 | T_+ [\mathcal{H}, T_-] | 0 \rangle}{\langle 0 | T_+ T_- | 0 \rangle}, \quad (1)$$

where $|0\rangle$ is the ground state wave function of the nucleus and T_{\pm} are the raising and lowering isospin operators. Hence, it is now clear that terms in the Hamiltonian not commuting with the isospin operator will contribute to E_{IAS} . In the early literature it was already shown that only the Coulomb interaction is not enough to explain available experimental data—this is known as the Nolen-Schiffer anomaly [20]—pointing to the fact that ISB of the nuclear strong interaction should play a non-negligible role.

In addition, it was shown that the energy of the IAS is intimately correlated with the neutron skin thickness of a heavy neutron-rich nucleus such as ^{208}Pb [21] and, thus, with the so-called symmetry energy [22,23]. The symmetry energy is a quantity very much studied in the context of neutron stars since it gives the penalty energy as a function of the density to convert all protons into neutrons in symmetric nuclear matter. Specifically, an approximate formula for E_{IAS} has been derived [21]. This formula explains in simple yet physical terms the relation of the E_{IAS} with the neutron skin thickness

$$E_{\text{IAS}} \approx \frac{6}{5} \sqrt{\frac{3}{5}} \frac{Ze^2}{r_p} \left(1 - \frac{1}{2} \frac{N}{N-Z} \frac{r_n - r_p}{r_p} \right), \quad (2)$$

where N and Z are the neutron and proton numbers, respectively; e the elementary electric charge; and r_p (r_n) the proton (neutron) root mean square radius. The neutron skin thickness is customarily defined as $r_n - r_p$. In EDFs r_p is commonly well constrained by experimental data on the electric charge radius of the nucleus while r_n or $r_n - r_p$ remains quite elusive and highly correlated with the neutron pressure inside nuclei or, which is the same, to the density dependence of the nuclear symmetry energy. The last simple equation predicts that the IAS energy should decrease with increasing neutron skin thickness. For example, for the case of ^{208}Pb Equation (2) is in reasonable agreement with the result of Hartree-Fock plus Random Phase approximation calculations presented in Ref. [21].

In this contribution, we will briefly analyze the possible ISB effects on the mass-radius relation of a neutron star by comparing two EDFs of the Skyrme type. The first one, SAMi [24], is of standard form and assumes exact isospin symmetry in the nuclear channel. The second one, SAMi-ISB [21], is identical to SAMi but contains all important ISB terms for a better description of the IAS energy. It is important to note that both functionals essentially use the same fitting protocol except for the fact that SAMi-ISB includes some extra information on ISB observables such as the E_{IAS} in ^{208}Pb and pseudo-observables such as the ISB contributions to the symmetric matter equation of state as calculated by using the realistic nucleon-nucleon AV18 interaction within the Brueckner-Hartree-Fock (BHF) approach [25].

The paper is organized as follows. In Section 2, we recall the basic form of ISB terms adopted here and although this has been published elsewhere [26], we deem it important to include it also here for the sake of completeness. In Section 3, we will show the main results of the present investigation while our conclusions are laid in Section 4.

2. Theory

We calculate the mass-radius relation of a cold non-accreting neutron star—neglecting the contribution from the crust—by solving the Tolman-Oppenheimer-Volkoff equation using as input neutron-proton-electron β -equilibrated matter. This procedure has been extensively used in the literature and further details can be found by the reader in Ref. [27].

To calculate the nuclear contribution assuming this model of neutron-star matter, we will use a nuclear Skyrme functional. Below we overview the main terms for our purposes here.

The isospin conserving form of the Skyrme functional adopted in this work can be easily found in the literature (see for example Equation (2.6) in Ref. [28]). Regarding the ISB contributions included in the SAMi-ISB functional [21] and neglected in the SAMi functional, we give a brief account here since those correspond to the main topic covered in this work. The contributions related to the Coulomb interaction discussed below will not contribute to the calculation of the neutron star mass-radius relation but definitely affect the fine tuning of the model in the fitting procedure that includes data on finite nuclei and, hence, will affect in an indirect way the results presented in Section 3.

2.1. Finite Size Effects

The electric charge distribution in momentum space is calculated by taking into account the effect of the proton and neutron electromagnetic form factors convoluted by using the proton and neutron point like densities, respectively, i.e., up to order $1/m^2$ where m stands for the mass of the nucleon and assuming spherical symmetry [29]

$$\begin{aligned} \rho_{\text{ch}}(q) &= \left(1 - \frac{q^2}{8m^2}\right) \left[G_{E,p}(q^2)\rho_p(q) + G_{E,n}(q^2)\rho_n(q)\right] \\ &- \frac{\pi q^2}{2m^2} \sum_{l,t} \left[2G_{M,t}(q^2) - G_{E,t}(q^2)\right] \langle \vec{l} \cdot \vec{s} \rangle \int_0^\infty dx \frac{j_1(qx)}{qx} |R_{n,l,j}(x)|^2 \end{aligned} \quad (3)$$

where $G_{E,M}$ are the electric (E) and magnetic (M) form factors [30], $R_{n,l,j}$ is the radial single particle wave function, t refers to the isospin state of the nucleon and n, l, j and s to the principal, orbital, total and spin quantum numbers. The sums run over all occupied states. The density calculated in such a way is then used for the calculation of the Coulomb energy functional. Finite size effects are frequently neglected and the proton density distribution is used instead in many EDF calculations of the Coulomb terms.

2.2. Electromagnetic Spin-Orbit

The electromagnetic spin-orbit correction to the nucleon single-particle (sp) energy can be written within a non-relativistic approximation as follows,

$$\Delta \varepsilon_i = \frac{\hbar^2}{2m^2 c^2} x_i \langle \vec{l}_i \cdot \vec{s}_i \rangle \int \frac{dr}{r} \frac{dU_C}{dr} u_i^2(r), \quad (4)$$

where i labels the i -th nucleon quantum numbers, m_i is the neutron or proton mass, U_C is the Coulomb energy potential—calculated at the Hartree-Fock level—and x_i is equal to $g_p - 1$ for protons and g_n for neutrons; $g_n = -3.82608545(90)$ and $g_p = 5.585694702(17)$ are the neutron and proton g -factors, respectively [31]. The total energy correction to the binding energy of a nucleus will be then $\Delta E_{\text{emso}} = \sum_{i=1}^A \Delta \varepsilon_i$. The effect of this term can be treated perturbatively since it produces a change in the nucleon single particle energies of tens of keV and, thus, negligible in the single particle wave functions.

As an exercise, assuming a uniformly charged sphere of radius R_C one can estimate the energy correction to the single particle energies [32],

$$\Delta \varepsilon_i \approx -\frac{\hbar^3}{2m_i^2 c} \frac{Z\alpha}{R_C^3} \langle \vec{l}_i \cdot \vec{s}_i \rangle x_i, \quad (5)$$

where α is the fine structure constant.

2.3. Coulomb Exchange

In connection with the Skyrme interaction, it is customary to adopt the so-called Slater approximation for the calculation of the Coulomb exchange. The mean-field corresponding to such an approximation can be written as,

$$U_C^{\text{Slater}} \varphi_i(\vec{r}) = -e^2 \left(\frac{3}{\pi} \right)^{1/3} \rho_p^{1/3} \varphi_i(\vec{r}), \quad (6)$$

where e is the elementary electric charge, ρ_p is the proton density and φ_i is the single particle wave function of the i -th nucleon. The exact exchange has instead the well-known non-local form,

$$U_C^{\text{exact}} \varphi_i(\vec{r}) = -e^2 \int d^3r' \frac{\varphi_j^*(\vec{r}') \varphi_j(\vec{r})}{|\vec{r} - \vec{r}'|} \varphi_i(\vec{r}'). \quad (7)$$

Since the early pioneering calculations done with the Skyrme force, the Slater approximation has been known to work quite well. In particular, the relative difference between the exact total energy and that calculated with the Slater approximation has been found to be of the order of 3–6% in Ref. [33]. As a consequence, single-particle energies undergo small changes of about hundreds of keV at most (see, e.g., Ref. [34] for numerical details).

2.4. Vacuum Polarization Correction

The lowest order correction in the fine-structure constant to the Coulomb potential eZ/r is estimated via the vacuum polarization correction. The vacuum polarization modifies the charge of the proton over a range of the order of the electron Compton wavelength. The virtual emission and absorption of an electron-positron pair gives rise to a repulsive potential V_{vp} that corrects the Coulomb repulsion of two protons. Taking into account the finite size of the nucleus, the correction potential can be written as follows [35],

$$V_{\text{vp}}(\vec{r}) = \frac{2}{3} \frac{\alpha e^2}{\pi} \int d^3r' \frac{\rho_{\text{ch}}(\vec{r}')}{|\vec{r} - \vec{r}'|} \mathcal{K}_1 \left(\frac{2}{\lambda_e} |\vec{r} - \vec{r}'| \right) \quad (8)$$

where e is the fundamental electric charge, α the fine-structure constant, λ_e the reduced Compton electron wavelength and

$$\mathcal{K}_1(x) \equiv \int_1^\infty dt e^{-xt} \left(\frac{1}{t^2} + \frac{1}{2t^4} \right) \sqrt{t^2 - 1}. \quad (9)$$

2.5. Charge Symmetry Breaking and Charge Independence Breaking Potentials

A charge symmetry breaking potential can be generally defined as $V_{\text{CSB}} = V_{nn} - V_{pp}$ while a charge independence breaking potential can be generally defined as $V_{\text{CIB}} = \frac{1}{2}(V_{nn} + V_{pp}) - V_{pn}$. So an effective nucleon-nucleon potential can be build on three parts: the charge independent part (isospin symmetry is fully preserved) plus the two previously defined breaking terms.

Following [36] (cf. Equations (18)–(21)) one can define Skyrme-like CSB and CIB potentials as follows,

$$V_{\text{CSB}}(\vec{r}_1, \vec{r}_2) \equiv \frac{1}{4} [\tau_z(1) + \tau_z(2)] \left\{ s_0(1 + y_0 P_\sigma) \delta(\vec{r}_1 - \vec{r}_2) + \frac{1}{2} s_1(1 + y_1 P_\sigma) \right. \\ \left. \times \left[P'^2 \delta(\vec{r}_1 - \vec{r}_2) + \delta(\vec{r}_1 - \vec{r}_2) P^2 \right] + s_2(1 + y_2 P_\sigma) \vec{P}' \cdot \delta(\vec{r}_1 - \vec{r}_2) \vec{P} \right\} \quad (10)$$

and

$$V_{\text{CIB}}(\vec{r}_1, \vec{r}_2) \equiv \frac{1}{2} \tau_z(1) \tau_z(2) \left\{ u_0(1 + z_0 P_\sigma) \delta(\vec{r}_1 - \vec{r}_2) + \frac{1}{2} u_1(1 + z_1 P_\sigma) \right. \\ \left. \times \left[P'^2 \delta(\vec{r}_1 - \vec{r}_2) + \delta(\vec{r}_1 - \vec{r}_2) P^2 \right] + u_2(1 + z_2 P_\sigma) \vec{P}' \cdot \delta(\vec{r}_1 - \vec{r}_2) \vec{P} \right\} \quad (11)$$

where $\vec{P} \equiv \frac{1}{2i}(\vec{\nabla}_1 - \vec{\nabla}_2)$ acts on the right, \vec{P}' is its complex conjugate acting on the left and $P_{\tau/\sigma}$ are the usual projector operators in isospin/spin spaces.

The Hartree-Fock energy density associated with the CSB and CIB potentials is (see appendix in Ref. [37])

$$\mathcal{H}_{\text{CSB}} = \frac{s_0}{8} (1 - y_0) (\rho_n^2 - \rho_p^2) \\ + \frac{3}{64} \{-s_1(1 - y_1) + s_2(1 + y_2)\} (\rho_n \nabla^2 \rho_n - \rho_p \nabla^2 \rho_p) \\ + \frac{1}{16} \{s_1(1 - y_1) + 3s_2(1 + y_2)\} (\rho_n \tau_n - \rho_p \tau_p) \\ + \frac{1}{32} \{s_1(1 - y_1) - s_2(1 + y_2)\} (J_n^2 - J_p^2) \quad (12a)$$

$$\mathcal{H}_{\text{CIB}} = \frac{u_0}{8} [(1 - z_0) (\rho_n^2 + \rho_p^2) - 2(2 + z_0) \rho_n \rho_p] \\ + \frac{3}{64} \{-u_1(1 - z_1) + u_2(1 + z_2)\} (\rho_n \nabla^2 \rho_n + \rho_p \nabla^2 \rho_p) \\ + \frac{1}{64} \{3u_1(2 + z_1) - u_2(2 + z_2)\} (\rho_n \nabla^2 \rho_p + \rho_p \nabla^2 \rho_n) \\ + \frac{1}{16} \{u_1(1 - z_1) + 3u_2(1 + z_2)\} (\rho_n \tau_n + \rho_p \tau_p) \\ - \frac{1}{16} \{u_1(2 + z_1) + u_2(2 + z_2)\} (\rho_p \tau_n + \rho_n \tau_p) \\ + \frac{1}{32} \{u_1(1 - 2z_1) - u_2(1 + 2z_2)\} (J_n^2 + J_p^2) \\ + \frac{1}{32} \{u_1 z_1 + u_2 z_2\} J^2 \quad (12b)$$

where τ , τ_n and τ_p are the total, neutron and proton kinetic densities, respectively, and J , J_n and J_p are the total, neutron and proton spin-orbit densities, respectively.

For the case of SAMi-ISB only the terms in s_0 and u_0 have been considered by setting y_0 and z_0 to -1 . The momentum-dependent terms were not considered since the information that we have at our disposal might not be sufficient to pin down the values of all parameters of a general interaction with several partial waves. We remind here that those terms set to zero in the Hamiltonian ((Equations (12a) and (12b)) will affect the neutron and proton effective masses as well as the central mean-field potential and spin-orbit potential [37]. Hence, we kept the model simple and with just a couple of extra parameters since the Coulomb corrections previously described are model independent. This is enough for SAMi-ISB to satisfactorily fit the IAS energy of ^{208}Pb and the symmetric matter EoS predicted by the AV18 potential when solved within the BHF approach.

The energy densities in Equations (12a) and (12b), needed for the calculation of the neutron star mass-relation, simplify as follows,

$$\mathcal{H}_{\text{CSB}} = \frac{s_0}{4} (\rho_n^2 - \rho_p^2) \quad \mathcal{H}_{\text{CIB}} = \frac{u_0}{4} [(\rho_n^2 + \rho_p^2) - \rho_n \rho_p], \quad (13)$$

being $u_0 = 25.8(4)$ MeV fm³ and $s_0 = -26.3(7)$ MeV fm³ determined from nuclear data and pseudo-data as previously explained. Summing and re-writing the previous terms,

$$\begin{aligned}\mathcal{H}_{\text{ISB}} \equiv \mathcal{H}_{\text{CIB}} + \mathcal{H}_{\text{CSB}} &= \frac{u_0 + s_0}{4} \rho_n^2 + \frac{u_0 - s_0}{4} \rho_p^2 + \frac{u_0}{4} \rho_n \rho_p \\ &= \frac{s_0}{4} \rho^2 \delta + \frac{u_0}{16} \rho^2 (1 + 3\delta^2),\end{aligned}\quad (14)$$

where $\delta \equiv (\rho_n - \rho_p)/\rho$ and $\rho = \rho_n + \rho_p$. Neutron star matter is characterized by values of $\delta \rightarrow 1$, and therefore,

$$\begin{aligned}\mathcal{H}_{\text{ISB}} &= \frac{s_0}{4} \rho^2 \delta + \frac{u_0}{16} \rho^2 [4 - 6(1 - \delta) + 3(1 - \delta)^2] \\ &\approx -\frac{u_0}{8} \rho^2 + \frac{3u_0 + 2s_0}{8} \rho^2 \delta \approx -\frac{u_0}{8} \rho^2 (1 - \delta),\end{aligned}\quad (15)$$

if u_0 were exactly equal to $-s_0$, a situation allowed from the statistical errors reported in the determination of the parameters. It is also interesting to note that for pure neutron matter ($\delta = 1$), the resulting energy density will be,

$$\mathcal{H}_{\text{ISB}} = \frac{u_0 + s_0}{4} \rho_n^2 \approx -\frac{1}{8} \rho_n^2 \text{ MeV fm}^{-3}, \quad (16)$$

with a statistical error in the determination of these parameters in the fitting protocol of SAMi-ISB that would be compatible with no ISB contribution to the neutron matter equation of state (as seen already in Equation (15)). Hence, it is important to note the relevance of the fine tuning of the parameters since the cancellation for this phenomenological choice of the CSB and CIB terms seems to be very large. The resulting energy contribution is negative and small, indicating a slight attractive overall behavior in pure neutron matter of ISB terms. The sign of this result is consistent with the neutron-neutron scattering length in the vacuum (cf. Table 1 in Ref. [38]). Since \mathcal{H}_{ISB} is included in the fitting protocol of the self-consistent SAMi-ISB functional, the isospin conserving part of the functional will be slightly modified with respect to that of SAMi in order to accommodate an accurate reproduction of the nuclear data (cf. Table 1). The effect on the observables such as the neutron matter EoS of such refit of the parameters can overcome, and will be shown to overcome, the genuine ISB effects. To which extent this cancellation effect depends on the uncertainties associated with the EDF fitting procedure, is an interesting subject for further investigation.

3. Results

In this section, we first compare the results of SAMi and SAMi-ISB for some ground and excited state properties of different nuclei as well as some relevant properties of the nuclear equation of state.

In Table 1, the binding energies and charge radii of some selected neutron-rich nuclei are given. The accuracy of both models is very good for the medium-heavy and heavy nuclei shown and never worse than a few ‰ for binding energies and 1% for charge radii. In the same table, some key nuclear matter properties at saturation density are given. These values, although non-observable, are thought to be in a reasonable range to properly describe not only binding energies and charge radii as briefly discussed above but also collective excited states such as the Giant Monopole, Dipole or Quadrupole Resonances (cf. Table 2 in Ref. [16]).

Table 1. Some selected neutron-rich nuclei observables and nuclear matter properties of the SAMi [24] and SAMi-ISB [21] functionals. Quantities within parenthesis correspond to the statistical theoretical error.

Observable	Nucleus	SAMi	SAMi-ISB	Experiment	Property	SAMi	SAMi-ISB
B (MeV)	^{48}Ca	415.61	417.67	415.99	ρ_0 (fm^{-3})	0.159(1)	0.1613(6)
	^{90}Zr	781.26	783.60	783.89	e_0 (MeV)	−15.93(9)	−16.03(2)
	^{132}Sn	1103.09	1102.75	1102.85	m_{IS}^*/m	0.6752(3)	0.730(19)
	^{208}Pb	1636.61	1635.78	1636.43	m_{IV}^*/m	0.664(13)	0.667(120)
					J (MeV)	28(1)	30.8(4)
					L (MeV)	44(7)	50(4)
					K_0 (MeV)	245(1)	235(4)
$\langle r_{\text{ch}}^2 \rangle^{1/2}$ (fm)	^{48}Ca	3.51	3.49	3.47			
	^{90}Zr	4.27	4.26	4.27			
	^{132}Sn	4.73	4.73	–			
	^{208}Pb	5.50	5.50	5.50			

In Figure 1, left panel, we show ISB contributions predicted by a Brueckner-Hartree-Fock (BHF) calculation using the AV18 realistic nucleon-nucleon potential to the energy per particle of symmetric nuclear matter as a function of the density [25]. We also show in the same panel the result from the SAMi-ISB functional that has been fitted to these pseudo-data.

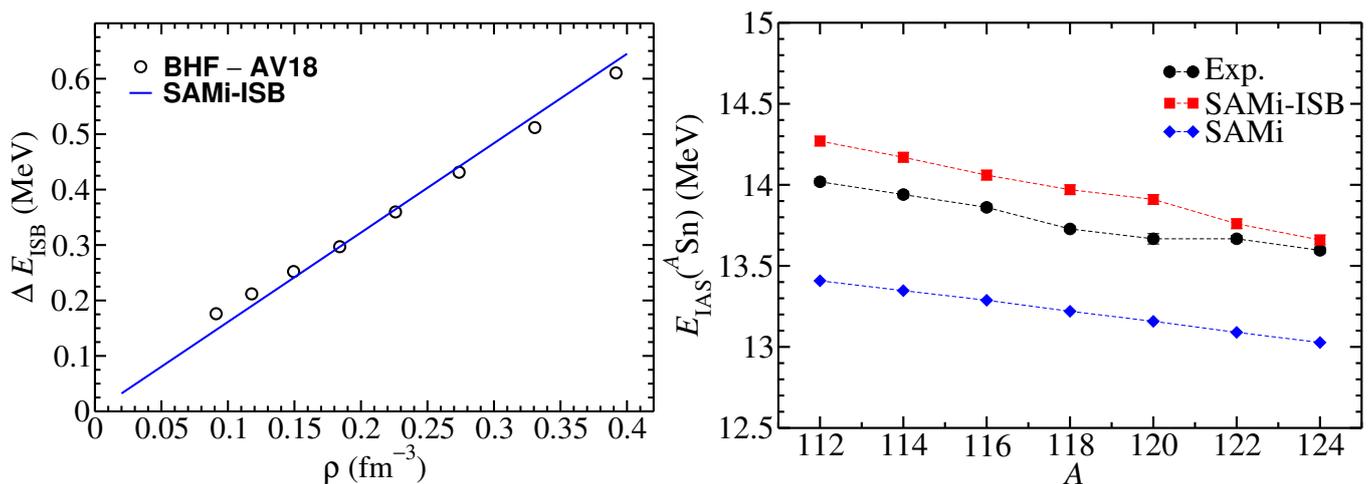


Figure 1. (Left panel) ISB contributions to the symmetric matter equation of state as predicted by the BHF calculations based on AV18 potential (black circles) of Ref. [25]. The solid line corresponds to SAMi-ISB. (Right panel) IAS energy for Sn isotopes as predicted by SAMi and SAMi-ISB, compared to experimental data. Figure taken from the supplemental material of Ref. [21].

As discussed in the introduction, the energy of the IAS is particularly sensitive to ISB effects. In Figure 1, right panel, we show the IAS energy for Sn isotopes as predicted by SAMi and SAMi-ISB, compared to experimental data. These results correspond to a prediction of the model since SAMi-ISB has been fitted to the IAS energy in ^{208}Pb alone. From that figure it is clear that ISB terms help in reproducing the absolute value of IAS energies along the Sn isotopic chain.

In Figure 2, left panel, we compare the neutron matter and symmetric matter EoS as predicted by SAMi and SAMi-ISB. Symmetric matter EoS are very similar from the low density region to $2\rho_0$ where they start to separate, with SAMi-ISB giving a softer increase of the energy per particle with density. The pure neutron matter EoS predicted by SAMi-ISB is slightly larger for densities up to about $\sim 0.5 \text{ fm}^{-3}$, where it becomes smaller, but still its trend is softer than SAMi for most of the density range shown. This behaviour is due to the fine-tuning of the SAMi-ISB parameters since ISB terms alone would produce attraction in neutron matter (cf. Equation (16)) and, thus, the opposite trend for $\rho \leq 0.5 \text{ fm}^{-3}$. Hence, the small ISB contributions cannot be disentangled from the

effect resulting from the re-fit of the parameters between SAMi and SAMi-ISB (cf. previous Section). In the inset, the same neutron matter EoSs are shown and compared with the results of two representative ab initio approaches for the low density part. The details of those calculations can be found in Refs. [39,40]. It is seen from the figure that SAMi-ISB agrees better than SAMi with the ab initio calculations shown here.

In the right panel of the same figure, the mass-radius relation of a neutron star is shown for the same models. The maximum observed mass of the J1614-2230 pulsar ($1.97 \pm 0.04 M_{\text{sun}}$) is also shown (grey horizontal band) [41]. In Ref. [42] an observation of the pulsar J0348 + 0432 allowed to estimate its mass as $2.01 \pm 0.04 M_{\text{sun}}$ which essentially sets the same maximum mass constraint (not shown in the figure). For completeness, we also show the simultaneous estimation (with 1σ error bars) of the mass and radius of the millisecond pulsar PSR J0030 + 0451 [43–45]. It is seen that both neutron star EoSs predict a radius that would be compatible with this observation at the 2σ level and that the SAMi-ISB model does not reproduce a maximum mass of $2M_{\text{sun}}$. On the other side, it is worth noting that both models are fully consistent with GW170817 [46] and GW190425 [47] radii constraints. In addition, the inclusion of a realistic crust is expected to increase the predicted radius for low and, to a lesser extent, medium mass stars while keeping the maximum mass prediction much less affected (cf. right panel of Figure 2.1 in Ref. [7]) and, thus, both models would likely be in better agreement with the radius observations of the PSR J0030+0451 pulsar. The brown shaded area encloses SAMi-ISB predictions with the predictions of the same model if $u_0 = s_0 = 0$. Those could be regarded as pure ISB effects while the difference between SAMi and SAMi-ISB is also affected by the result of the fitting procedure. Hence, even though large systematic uncertainties exist, it is clear from the figure that ISB effects that are known to be small in nuclei, may entail non-negligible effects on observables that are sensitive to large densities as those associated with the core of a neutron star. These results should be confirmed and tested on the bases of other models for neutron star matter.

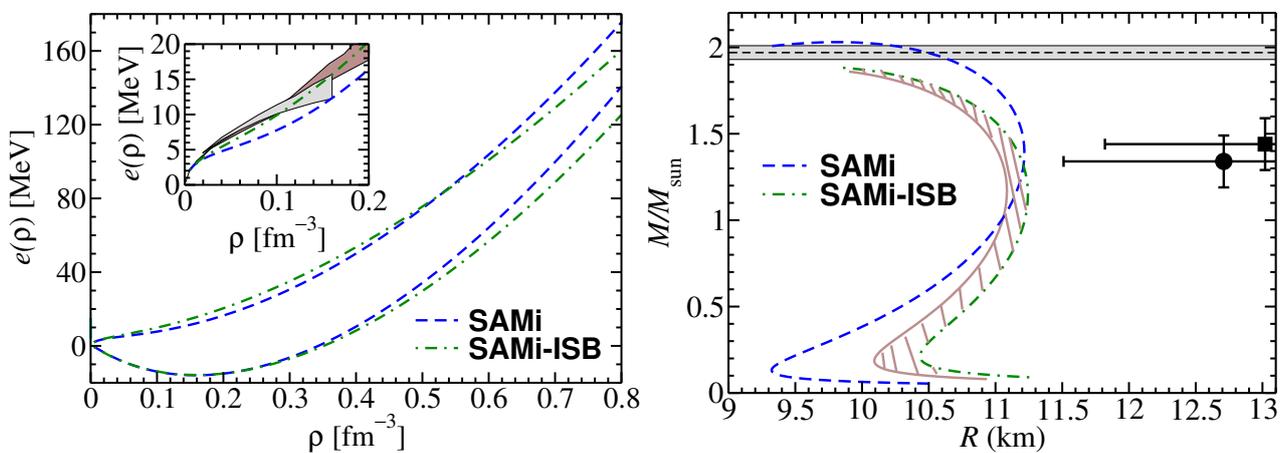


Figure 2. (Left panel) EoS of pure neutron and symmetric nuclear matter as predicted by SAMi (dashed) and SAMi-ISB (dash-dotted). In the inset only the neutron matter EoSs are shown and compared with the results from the ab initio approaches of Ref. [39] (brown shaded area) and Ref. [40] (grey shaded area). (Right panel) Mass-Radius relation of a neutron star as predicted by SAMi and SAMi-ISB. The brown shaded area encloses SAMi-ISB predictions with the predictions of the same model if $u_0 = s_0 = 0$. The maximum observed mass of the millisecond pulsar J1614-2230 ($1.97 \pm 0.04 M_{\text{sun}}$) is also shown (grey horizontal band) [41] as well as the simultaneous estimation of the mass and radius of the millisecond pulsar PSR J0030+0451 [43] (dot) and [44] (square). Errors correspond to the 1σ estimation.

For completeness, in Table 2, we report on some basic properties of a neutron star as predicted by SAMi, SAMi-ISB and SAMi-ISB with $u_0 = s_0 = 0$ functionals. Specifically, the maximum mass M_{max} and radius R_{max} where for SAMi-ISB the maximum mass is reduced by $\sim 7\%$ with respect to the one predicted by SAMi. The central density $\rho_{1.4}^c$ of a $1.4M_{\text{sun}}$ neutron star is $\sim 10\%$ larger for SAMi-ISB and the corresponding radius has

only slightly decreased. The sensitivity of the radius to nuclear ISB effects (see table and compare green dash-dotted and brown solid curves in the left panel of Figure 2) seems to be around hundreds of meters for a wide range of neutron star masses. The tidal deformability $\Lambda_{1.4}$ (see, e.g., [46,48–50]) predicted by SAMi-ISB shows a decrease with respect to SAMi that is larger than $\sim 10\%$. The range deduced from GW170817 is $\Lambda_{1.4} \leq 800$ [46]. The tidal deformability Λ is customarily written as a function of the compactness parameter $\zeta = GM/Rc^2$ (also shown in Table 2) and of the quadrupole tidal Love number k_2 [51] as $\Lambda = 2k_2\zeta^{-5}/3$.

In summary, according to SAMi-ISB results, we shall expect an influence of ISB terms originated from the nuclear strong interaction for properties related to very high densities—easily found in a neutron star interior—as well as a small to moderate effect on the radius and tidal deformability. The properties of the crust, not studied here, might also be affected by ISB effects whenever competing configurations are close enough in energy since ISB effects at normal densities are expected to be of the order of hundreds of keV.

Table 2. Some basic properties of a neutron star as predicted by SAMi, SAMi-ISB and SAMi-ISB with $u_0 = s_0 = 0$ functionals. Specifically, the maximum mass (M_{\max}), radius (R_{\max}) and central density ($\rho_{1.4}^c$), radius ($R_{1.4}$), tidal deformability ($\Lambda_{1.4}$) and compactness ($\zeta_{1.4}$) for a $1.4M_{\text{sun}}$ neutron star.

	M_{\max}/M_{sun}	R_{\max} [km]	$\rho_{1.4}^c$ [fm^{-3}]	$R_{1.4}$	$\Lambda_{1.4}$ [km]	$\zeta_{1.4}$
SAMi	2.03	9.8	0.54	11.2	301	0.18
SAMi-ISB	1.88	9.8	0.59	11.2	261	0.19
SAMi-ISB ($u_0 = s_0 = 0$)	1.86	9.9	0.61	11.0	242	0.19

4. Conclusions

In this contribution, we present a preliminary study of the effect of ISB on the mass-radius relation of a neutron star and a few related observables such as the tidal deformability. Since densities expected along the core of the star are thought to be well above $2\rho_0$, ISB terms that are commonly small and frequently negligible for the study of finite nuclear properties, may become relevant for the study of neutron star matter and specifically its mass-radius relation (see brown area in the right panel of Figure 2). We compared the predictions of two models based on the same fitting protocol, one containing non-standard ISB terms (SAMi-ISB) and the other neglecting them (SAMi). Some differences appear even though there is a large cancellation between the CSB and CIB parts of the proposed phenomenological interaction. The extrapolation of our model to very large densities should be further investigated and complemented with other type of calculations in order to confirm or dismiss the effects of ISB on the mass-radius relation of a neutron star and, whenever possible, study the large cancellation predicted here for the CSB and CIB terms. The study of ISB effects on other observable properties of neutron stars might also be envisaged.

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