

Effect of beam energy straggling on resonant yield in thin gas targets: The cases $^{22}\text{Ne}(p, \gamma)^{23}\text{Na}$ and $^{14}\text{N}(p, \gamma)^{15}\text{O}$

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Abstract – When deriving resonance strengths using the thick-target yield approximation, for very narrow resonances it may be necessary to take beam energy straggling into account. This applies to gas targets of a few keV width, especially if there is some additional structure in target stoichiometry or detection efficiency. The correction for this effect is shown and tested on recent studies of narrow resonances in the $^{22}\text{Ne}(p, \gamma)^{23}\text{Na}$ and $^{14}\text{N}(p, \gamma)^{15}\text{O}$ reactions.

1 Introduction. – In the thick target study of a radiative proton capture reaction, the yield on top of the resonance plateau Y_{max} is used to obtain the resonance strength $\omega\gamma$ as:

$$\omega\gamma = \omega \frac{\Gamma_p \Gamma_\gamma}{\Gamma_p + \Gamma_\gamma} = \frac{2 Y_{\text{max}} \epsilon_R}{\lambda_R^2} \frac{m_t}{m_t + m_p} \quad (1)$$

where ω is the statistical factor, $\Gamma_{p,\gamma}$ are the proton and γ -ray widths of the resonance under study, ϵ_R is the effective stopping power in the laboratory system, λ_R^2 is the

squared de Broglie wavelength at the center-of-mass resonance energy, and m_t and m_p are the masses of target and projectile, respectively. Equation (1) is applicable when the energetic target thickness ΔE is large compared to the total width $\Gamma = \Gamma_p + \Gamma_\gamma$. For intermediate cases, the strength from eq. (1) must be multiplied with a factor $\frac{\pi}{2} / \arctan(\Delta E / \Gamma)$ that approaches unity for $\Delta E \gg \Gamma$ [1].

For the experiment of Refs. [2, 3], $\Delta E \sim 3.9$ keV (Figure 2). The three new resonances reported in Refs. [2, 3] correspond to excited states in ^{23}Na at $E_x = 8944$ ($J^\pi = 3/2^+$), 8975 ($5/2^+$), and 9042 ($7/2^+$ or $9/2^+$) keV, respectively

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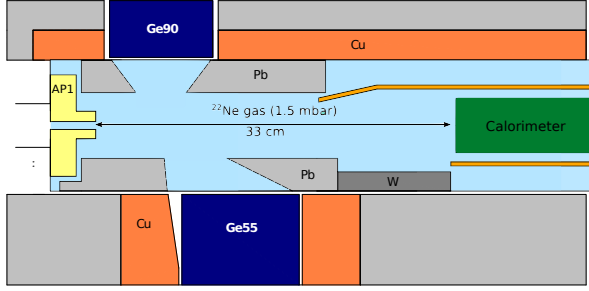


Fig. 1: Sketch of the experimental setup. The ion beam enters from the left and is stopped on the beam calorimeter. The color encodes the HPGe detectors (Ge55 and Ge90, dark blue), the copper (Cu, orange), lead (Pb, light grey), and tungsten (W, dark grey) shielding, the gas limiting aperture (AP1, yellow), the beam calorimeter (green), and the ^{22}Ne gas (light blue).

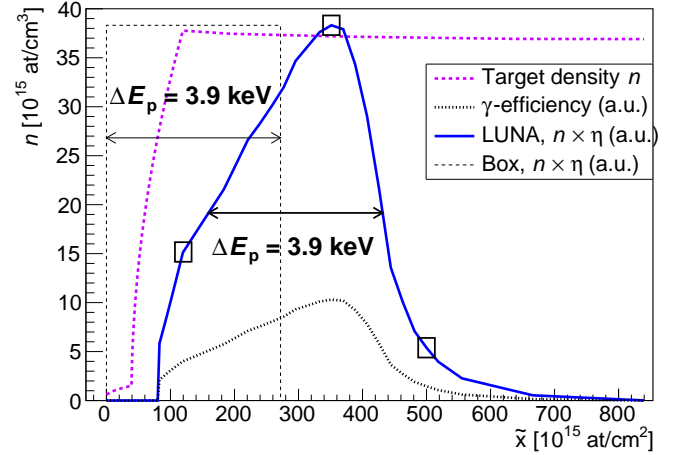


Fig. 2: LUNA gas target density (n , purple dashed line), γ -detection efficiency (η , black dotted line, for $E_\gamma = 478$ keV, detector Ge55 [2, 3]), and product of the two ($n \times \eta$, blue straight line) as a function of position \tilde{x} in the target. An ideal box shape of same width and maximum is also shown (black dashed line). See text for details.

(spin assignments from Ref. [4]). They were not observed in a nuclear resonance fluorescence experiment with bremsstrahlung up to 10.4 MeV and a typical sensitivity of 0.1 eV for the partial γ -width [5]. Assuming 1% ground state branching, this leads to a limit of $\Gamma_\gamma < 10$ eV. In a $^{22}\text{Ne}(^3\text{He},d)^{23}\text{Na}$ experiment [6], values or upper limits between 2.3×10^{-9} eV and 1.1×10^{-6} eV are reported for the proton widths of these states.

As a result, $\Delta E \gg \Gamma_\gamma + \Gamma_p = \Gamma$ is found for all three resonances, and they seem to be textbook [1] cases for the applicability of the thick-target yield formula eq. (1). This is the assumption made explicitly in Refs. [2, 3]. However, this is inappropriate, as will be shown below.

Derivation of the correction. – The experimental setup (Figure 1) is characterized by a windowless, static-type ^{22}Ne gas target with two HPGe detectors, collimated to give effective detection angles of 55° , and 90° , respectively.

The profiles of effective target density n and γ -ray detection efficiency η for the present case [2, 3, 7, 8] are shown in Figure 2. For their product $n \times \eta$, a full width at half maximum of $\Delta E_p = 3.9$ keV is found. There is significant structure in the γ -ray detection efficiency curve as a consequence of the collimation. For the sake of the discussion, an ideal box profile of the same width and with a height corresponding to the maximum of the LUNA $n \times \eta$ profile has been added (Figure 2). This profile starts at the same place as the real gas target, i.e. in the connection tube between final collimator and first pumping stage.

As a next step, the energy loss of an $E_p = 190$ keV proton beam when passing through the ^{22}Ne gas target is simulated by SRIM [9]. The energy distribution of the slowed beam is almost Gaussian (Figure 3) with mean en-

ergy $E_{\text{slowed}}^{\text{mean}}$ and straggling width σ_{strag} , and an empirical parameterization for $\sigma_{\text{strag}}(E_{\text{slowed}}^{\text{mean}})$ is derived.

Using this information and in addition the energy spread of the proton beam from the accelerator (at LUNA, $\sigma_{\text{beam}} \leq 0.1/2.355$ keV [10], much lower than the straggling width), an approximate energy distribution of the slowed beam at position \tilde{x} in the target can be derived as:

$$f_{\text{beam}}(E, E_{\text{slowed}}^{\text{mean}}) = \exp \left[-\frac{(E - E_{\text{slowed}}^{\text{mean}})^2}{2\sigma_{\text{strag}}^2(E_{\text{slowed}}^{\text{mean}}) + 2\sigma_{\text{beam}}^2} \right] \quad (2)$$

Finally, the yield $Y(E_p)$ for a resonance scan performed on this target with this beam can be computed for each incident beam energy E_p , by numerically integrating the product of the usual Breit-Wigner resonance shape $\sigma_{\text{BW}}(E)$, $f_{\text{beam}}(E, E_{\text{slowed}}^{\text{mean}})$, and the target and efficiency profiles $n(\tilde{x}) \times \eta(\tilde{x})$:

$$Y(E_p) = \int_{\tilde{x}=0}^{\tilde{x}=\tilde{x}_{\text{max}}} d\tilde{x} \int_{E=E_p}^{E=0} dE \sigma_{\text{BW}}(E) f_{\text{beam}}(E, E_{\text{slowed}}^{\text{mean}}(\tilde{x})) n(\tilde{x}) \eta(\tilde{x}) \quad (3)$$

As a first step, this formula is applied to the ideal box profile for the new $E_p = 189.5$ keV resonance, first by artificially imposing $\sigma_{\text{strag}} \equiv 0$ so that only the small beam energy spread remains (black thin dashed curve in Figure 4), then including σ_{strag} from SRIM (black thin straight curve in Figure 4). It is thus confirmed that in the textbook case, straggling does not significantly affect the plateau yield [1].

However, the picture changes when the realistic profile (blue curve in Figure 2) is used: Here, the artificial case without straggling (thick blue dashed curve in Figure 4) is

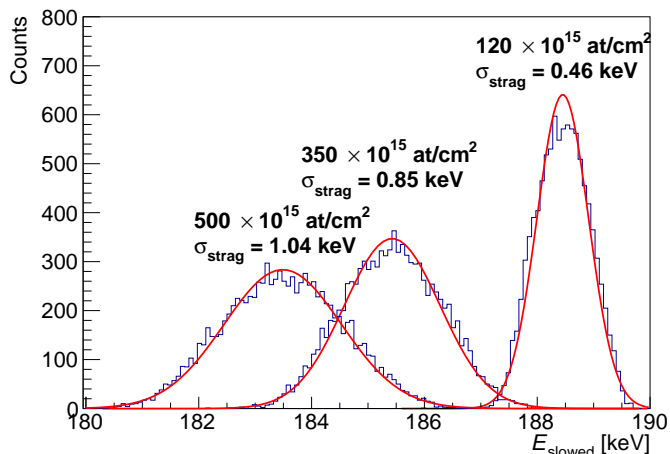


Fig. 3: Energy straggling of the slowed proton beam, simulated by SRIM [9], at the three points marked in Figure 2.

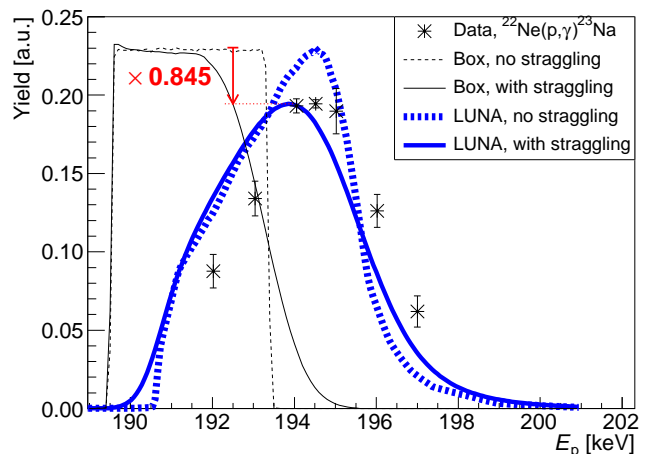


Fig. 4: Experimental yield (data points) of the $E_\gamma = 440$ keV transition in the 189.5 keV $^{22}\text{Ne}(p,\gamma)^{23}\text{Na}$ resonance, compared with calculated yield curves without straggling (blue dotted line) and with straggling (blue full line). The experimental yield is about 15% lower than expected for an ideal flat-top profile without straggling (thin dashed black line), or with straggling (thin black line). See text for details.

76 significantly higher than the realistic case with straggling
77 (thick blue full curve in Figure 4).

78 As a consequence, the resonance strength data by
79 Refs. [2, 3] must be corrected upwards, dividing by the
80 factor C that reflects the reduction in the yield profile
81 observed when correctly including the energy straggling
82 of the proton beam. Repeating the calculation also for
83 the other two resonances, new values for the resonance
84 strength are found (Table 1). Since this correction is a
85 purely calculated one, a conservative error bar of 50% is
86 assigned to the correction.

87 The present, corrected resonance strengths are in agree-
88 ment with independent, new data on the resonances at
89 156.2 and 189.5 keV that have been reported in the mean
90 time by the TUNL group [11]. The upper limits reported
91 in Refs. [2, 3] are not updated here, because much more
92 restrictive upper limits may be expected from an experi-
93 ment on the same reaction with a new setup [12].

94 **Application to the $^{14}\text{N}(p,\gamma)^{15}\text{O}$ reaction.** – In
95 order to gain a further cross-check of the validity of the
96 present correction, a run of this measurement campaign
97 [2, 3] is re-analyzed here. During this run aimed at ex-
98 tending the γ -detection efficiency curve to higher energies,
99 the gas target was filled with 3.5 mbar nitrogen, and the
100 $E_p = 278$ keV resonance in the $^{14}\text{N}(p,\gamma)^{15}\text{O}$ reaction was
101 excited. It is well known that the γ rays de-exciting this
102 $1/2^+$ resonance are isotropic [13, 14], and the branching
103 for the strongest branch, consisting of two γ rays at 1384
104 and 6172 keV, respectively, has recently been remeasured
105 very precisely, to $(58.3 \pm 0.3)\%$ [15–17].

106 This resonance has a non-negligible total width, $\Gamma_p =$
107 (0.999 ± 0.046) keV [17] that must be included in eq. (2).
108 Even though the target is much wider than the resonance,
109 $\Delta E_p \sim 19.4$ keV, again due to the structure of the efficiency
110 curve, a correction is necessary, by dividing by 0.890. Us-
111 ing the known branching [15–17], a resonance strength of
112 $(12.7 \pm 0.3_{\text{stat}} \pm 1.0_{\text{syst}})$ meV is found. The systematic er-

ror includes 4% for the γ -detection efficiency, 1% for the
beam intensity, 2.9% for the stopping of protons in nitro-
gen [9], 0.5% for the branching ratio, and 6% for the
present correction, in total 8%.

117 The present new strength for the 278 keV resonance
118 in the $^{14}\text{N}(p,\gamma)^{15}\text{O}$ reaction is consistent with a recent
119 very precise re-evaluation, which resulted in a value of
120 (12.6 ± 0.3) meV [17].

121 **Discussion.** – The presently derived correction for
122 narrow resonance yields, while general in nature, is ex-
123 pected to play a significant role only for the case of tar-
124 gets that are extended over such a large space that in-
125 homogeneities in both target thickness and detection effi-
126 ciency must be considered. In practice this will mainly ap-
127 ply to windowless, static-type gas targets, whose density-
128 efficiency profile, $n(x) \times \eta(x)$, significantly deviates from
129 an ideal box shape.

130 Of all the LUNA experiments, this correction has a sig-
131 nificant effect only for the HPGe-detector based phase of
132 the study of the $^{22}\text{Ne}(p,\gamma)^{23}\text{Na}$ reaction [2, 3, 8]. In the
133 other LUNA experiments [18], either non-resonant pro-
134 cesses were studied [19–23, e.g.] or solid targets (where
135 the detection efficiency inside the target is approximately
136 constant and structure usually mitigated) were used [24–
137 26, e.g.].

138 It should be noted that in these cases, the alternative
139 approach to integrate the entire resonance scan [1] cannot
140 be applied due to the non-constant detection efficiency.

141 A static-type gas target should be designed in such
142 a way that not only the total width of the resonance,
143 but also the beam energy spread (loss and straggling) re-
144 main much smaller than any $n(\tilde{x}) \times \eta(\tilde{x})$ structure. This

Table 1: Original [2, 3] and corrected [8] values for the resonance strength $\omega\gamma$, and straggling correction factor C . For the corrected values, the statistical error bars (unchanged) and the systematical error bars (including the new correction) are given separately.

E_p [keV]	$\omega\gamma_{\text{orig}}$ [eV] [2, 3]	C	$\omega\gamma_{\text{corr}}$ [eV] [8]	stat.	syst.
156.2	$(1.48 \pm 0.10) \times 10^{-7}$	0.845	1.8×10^{-7}	6%	8%
189.5	$(1.87 \pm 0.06) \times 10^{-6}$	0.850	2.2×10^{-6}	2%	8%
259.7	$(6.89 \pm 0.16) \times 10^{-6}$	0.841	8.2×10^{-6}	1%	8%

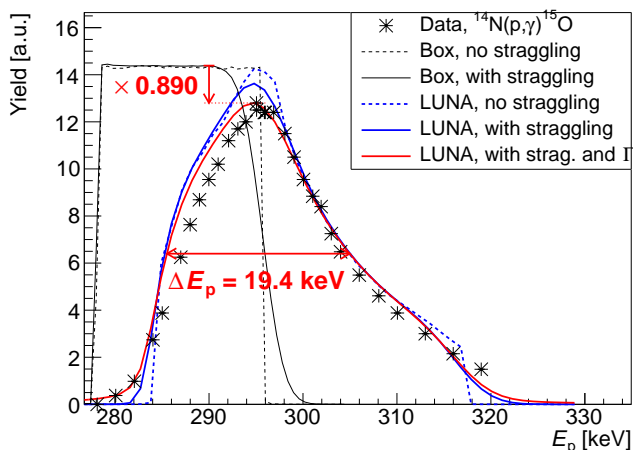


Fig. 5: Same as Figure 4, but for the 278 keV resonance in the $^{14}\text{N}(p,\gamma)^{15}\text{O}$ reaction. The final correction includes not only the straggling, but also a total resonance width of 1.07 keV in the laboratory system.

is actually the case in forthcoming LUNA work on the $^{22}\text{Ne}(p,\gamma)^{23}\text{Na}$ reaction with higher target density and very flat efficiency profile, due to the use of a γ -calorimeter [12].

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*. – REFERENCES

- [1] ILIADIS C., *Nuclear Physics of Stars* 2nd Edition (Wiley-VCH, Weinheim) 2015.
- [2] CAVANNA F., DEPALO R., ALIOTTA M., ANDERS M., BEMMERER D., BEST A., BOELTZIG A., BROGGINI C., BRUNO C. G., CACIOLLI A., CORVISIERO P., DAVINSON T., DI LEVA A., ELEKES Z., FERRARO F., FORMICOLA A., FÜLÖP Z., GERVINO G., GUGLIELMETTI A., GUSTAVINO C., GYÜRKY G., IMBRIANI G., JUNKER M., MENEGAZZO R., MOSSA V., PANTALEO F. R., PRATI P., SCOTT D. A., SOMORJAI E., STRANIERO O., STRIEDER F., SZÜCS T., TAKÁCS M. P., TREZZI D. LUNA COLLABORATION, *Phys. Rev. Lett.*, **115** (2015) 252501.

- [3] DEPALO R., CAVANNA F., ALIOTTA M., ANDERS M., BEMMERER D., BEST A., BOELTZIG A., BROGGINI C., BRUNO C. G., CACIOLLI A., CIANI G. F., CORVISIERO P., DAVINSON T., DI LEVA A., ELEKES Z., FERRARO F., FORMICOLA A., FÜLÖP Z., GERVINO G., GUGLIELMETTI A., GUSTAVINO C., GYÜRKY G., IMBRIANI G., JUNKER M., MENEGAZZO R., MOSSA V., PANTALEO F. R., PIATTI D., PRATI P., STRANIERO O., SZÜCS T., TAKÁCS M. P., TREZZI D., *Phys. Rev. C*, **94** (2016) 055804.
<http://link.aps.org/doi/10.1103/PhysRevC.94.055804>

- [4] JENKINS D., BOUHELAL M., COURTIN S., FREER M., FULTON B. *et al.*, *Phys. Rev. C*, **87** (2013) 064301.

- [5] VODHANEL R., BRUSSEL M. K., MOREH R., SELLYEY W. C., CHAPURAN T. E., *Phys. Rev. C*, **29** (1984) 409.

- [6] HALE S. E., CHAMPAGNE A. E., ILIADIS C., HANSPER V. Y., POWELL D. C., BLACKMON J. C., *Phys. Rev. C*, **65** (2001) 015801.

- [7] CAVANNA F., DEPALO R., MENZEL M.-L., ALIOTTA M., ANDERS M., BEMMERER D., BROGGINI C., BRUNO C. G., CACIOLLI A., CORVISIERO P., DAVINSON T., DI LEVA A., ELEKES Z., FERRARO F., FORMICOLA A., FÜLÖP Z., GERVINO G., GUGLIELMETTI A., GUSTAVINO C., GYÜRKY G., IMBRIANI G., JUNKER M., MENEGAZZO R., PRATI P., ROSSI ALVAREZ C., SCOTT D. A., SOMORJAI E., STRANIERO O., STRIEDER F., SZÜCS T., TREZZI D., *Eur. Phys. J. A*, **50** (2014) 179.

- [8] CAVANNA F., DEPALO R., ALIOTTA M., ANDERS M., BEMMERER D., BEST A., BOELTZIG A., BROGGINI C., BRUNO C. G., CACIOLLI A., CORVISIERO P., DAVINSON T., DI LEVA A., ELEKES Z., FERRARO F., FORMICOLA A., FÜLÖP Z., GERVINO G., GUGLIELMETTI A., GUSTAVINO C., GYÜRKY G., IMBRIANI G., JUNKER M., MENEGAZZO R., MOSSA V., PANTALEO F. R., PRATI P., SCOTT D. A., SOMORJAI E., STRANIERO O., STRIEDER F., SZÜCS T., TAKÁCS M. P., TREZZI D. LUNA COLLABORATION, *Phys. Rev. Lett.*, **120** (2018) 239901(E).

- [9] ZIEGLER J. F., ZIEGLER M. D. BIERSACK J. P., *Nucl. Inst. Meth. B*, **268** (2010) 1818.
<http://dx.doi.org/10.1016/j.nimb.2010.02.091>
- [10] FORMICOLA A., IMBRIANI G., JUNKER M., BEMMERER D., BONETTI R., BROGGINI C., CASELLA C., CORVISIERO P., COSTANTINI H., GERVINO G., GUSTAVINO C., LEMUT A., PRATI P., ROCA V., ROLFS C., ROMANO M., SCHÜRMANN D., STRIEDER F., TERRASI F., TRAUTVETTER H.-P. ZAVATARELLI S., *Nucl. Inst. Meth. A*, **507** (2003) 609.
- [11] KELLY K. J., CHAMPAGNE A. E., DOWNEN L. N., DERMIGNY J. R., HUNT S., ILIADIS C. COOPER A. L., *Phys. Rev. C*, **95** (2017) 015806.
- [12] FERRARO F., TAKÁCS M. P., PIATTI D., MOSSA V., ALIOTTA M., BEMMERER D., BEST A., BOELTZIG A., BROGGINI C., BRUNO C. G., CACIOLLI A., CAVANNA F., CHILLERY T., CIANI G. F., CORVISIERO P., CSEDREKI L., DAVINSON T., DEPALO R., D'ERASMO G., DI LEVA A., ELEKES Z., FIORE E. M., FORMICOLA A., FÜLÖP Z., GERVINO G., GUGLIELMETTI A., GUSTAVINO C., GYÜRKY G., IMBRIANI G., JUNKER M., KOCHANEK I., LUGARO M., MARCUCCI L. E., MARIGO P., MENEGAZZO R., PANTALEO F. R., PATICCHIO V., PERRINO R., PRATI P., SCHIAVULLI L., STÖCKEL K., STRANIERO O., SZÜCS T., TREZZI D. ZAVATARELLI S., *Eur. Phys. J. A*, **54** (2018) 44.
- [13] AJZENBERG-SELOVE F., *Nucl. Phys. A*, **523** (1991) 1.
- [14] COSTANTINI H., *Direct measurements of radiative capture reactions at astrophysical energies* Ph.D. thesis Università degli studi di Genova (2003).
- [15] MARTA M., FORMICOLA A., GYÜRKY G., BEMMERER D., BROGGINI C., CACIOLLI A., CORVISIERO P., COSTANTINI H., ELEKES Z., FÜLÖP Z., GERVINO G., GUGLIELMETTI A., GUSTAVINO C., IMBRIANI G., JUNKER M., KUNZ R., LEMUT A., LIMATA B., MAZZOCCHI C., MENEGAZZO R., PRATI P., ROCA V., ROLFS C., ROMANO M., ALVAREZ C. R., SOMORJAI E., STRANIERO O., STRIEDER F., TERRASI F., TRAUTVETTER H. P. VOMIERO A., *Phys. Rev. C*, **78** (2008) 022802(R).
- [16] MARTA M., FORMICOLA A., BEMMERER D., BROGGINI C., CACIOLLI A., CORVISIERO P., COSTANTINI H., ELEKES Z., FÜLÖP Z., GERVINO G., GUGLIELMETTI A., GUSTAVINO C., GYÜRKY G., IMBRIANI G., JUNKER M., LEMUT A., LIMATA B., MAZZOCCHI C., MENEGAZZO R., PRATI P., ROCA V., ROLFS C., ROSSI ALVAREZ C., SOMORJAI E., STRANIERO O., STRIEDER F., TERRASI F., TRAUTVETTER H. P. VOMIERO A., *Phys. Rev. C*, **83** (2011) 045804.
- [17] DAIGLE S., KELLY K. J., CHAMPAGNE A. E., BUCKNER M. Q., ILIADIS C. HOWARD C., *Phys. Rev. C*, **94** (2016) 025803.
- [18] BROGGINI C., BEMMERER D., CACIOLLI A. TREZZI D., *Prog. Part. Nucl. Phys.*, **98** (2018) 55.
<http://arxiv.org/abs/1707.07952>
- [19] FORMICOLA A., IMBRIANI G., COSTANTINI H., ANGULO C., BEMMERER D., BONETTI R., BROGGINI C., CORVISIERO P., CRUZ J., DESCOUVE-MONT P., FÜLÖP Z., GERVINO G., GUGLIELMETTI A., GUSTAVINO C., GYÜRKY G., JESUS A. P., JUNKER M., LEMUT A., MENEGAZZO R., PRATI P., ROCA V., ROLFS C., ROMANO M., ROSSI ALVAREZ C., SCHÜMANN F., SOMORJAI E., STRANIERO O., STRIEDER F., TERRASI F., TRAUTVETTER H. P., VOMIERO A. ZAVATARELLI S., *Phys. Lett. B*, **591** (2004) 61.
- [20] LEMUT A., BEMMERER D., CONFORTOLA F., BONETTI R., BROGGINI C., CORVISIERO P., COSTANTINI H., CRUZ J., FORMICOLA A., FÜLÖP Z., GERVINO G., GUGLIELMETTI A., GUSTAVINO C., GYÜRKY G., IMBRIANI G., JESUS A. P., JUNKER M., LIMATA B., MENEGAZZO R., PRATI P., ROCA V., ROGALLA D., ROLFS C., ROMANO M., ROSSI ALVAREZ C., SCHÜMANN F., SOMORJAI E., STRANIERO O., STRIEDER F., TERRASI F. TRAUTVETTER H. P., *Phys. Lett. B*, **634** (2006) 483.
- [21] BEMMERER D., CONFORTOLA F., COSTANTINI H., FORMICOLA A., GYÜRKY G., BONETTI R., BROGGINI C., CORVISIERO P., ELEKES Z., FÜLÖP Z., GERVINO G., GUGLIELMETTI A., GUSTAVINO C., IMBRIANI G., JUNKER M., LAUBENSTEIN M., LEMUT A., LIMATA B., LOZZA V., MARTA M., MENEGAZZO R., PRATI P., ROCA V., ROLFS C., ALVAREZ C. R., SOMORJAI E., STRANIERO O., STRIEDER F., TERRASI F. TRAUTVETTER H. P., *Phys. Rev. Lett.*, **97** (2006) 122502.
- [22] CACIOLLI A., MAZZOCCHI C., CAPOGROSSO V., BEMMERER D., BROGGINI C., CORVISIERO P., COSTANTINI H., ELEKES Z., FORMICOLA A., FÜLÖP Z., GERVINO G., GUGLIELMETTI A., GUSTAVINO C., GYÜRKY G., IMBRIANI G., JUNKER M., LEMUT A., MARTA M., MENEGAZZO R., PALMERINI S., PRATI P., ROCA V., ROLFS C., ROSSI ALVAREZ C., SOMORJAI E., STRANIERO O., STRIEDER F., TERRASI F., TRAUTVETTER H. P. VOMIERO A., *Astron. Astrophys.*, **533** (2011) A66.
- [23] ANDERS M., TREZZI D., MENEGAZZO R., ALIOTTA M., BELLINI A., BEMMERER D., BROGGINI C., CACIOLLI A., CORVISIERO P., COSTANTINI H., DAVINSON T., ELEKES Z., ERHARD M., FORMICOLA A.,

314 FÜLÖP Z., GERVINO G., GUGLIELMETTI A., GUS-
315 TAVINO C., GYÜRKY G., JUNKER M., LEMUT A.,
316 MARTA M., MAZZOCCHI C., PRATI P., ROSSI AL-
317 VAREZ C., SCOTT D. A., SOMORJAI E., STRANIERO
318 O. SZÜCS T., *Phys. Rev. Lett.*, **113** (2014) 042501.

319 [24] STRIEDER F., LIMATA B., FORMICOLA A., IM-
320 BRIANI G., JUNKER M., BEMMERER D., BEST
321 A., BROGGINI C., CACIOLLI A., CORVISIERO P.,
322 COSTANTINI H., DILEVA A., ELEKES Z., FÜLÖP
323 Z., GERVINO G., GUGLIELMETTI A., GUSTAVINO
324 C., GYÜRKY G., LEMUT A., MARTA M., MAZ-
325 ZOCCHI C., MENEGAZZO R., PRATI P., ROCA
326 V., ROLFS C., ROSSI ALVAREZ C., SOMORJAI E.,
327 STRANIERO O., TERRASI F. TRAUTVETTER H. P.,
328 *Phys. Lett. B*, **707** (2012) 60.

329 [25] SCOTT D. A., CACIOLLI A., DI LEVA A., FORMI-
330 COLA A., ALIOTTA M., ANDERS M., BEMMERER
331 D., BROGGINI C., CAMPEGGIO M., CORVISIERO
332 P., ELEKES Z., FÜLÖP Z., GERVINO G., GUGLIEL-
333 METTI A., GUSTAVINO C., GYÜRKY G., IMBRIANI
334 G., JUNKER M., LAUBENSTEIN M., MENEGAZZO
335 R., MARTA M., NAPOLITANI E., PRATI P., RIGATO
336 V., ROCA V., SOMORJAI E., SALVO C., STRANIERO
337 O., STRIEDER F., SZÜCS T., TERRASI F. TREZZI
338 D., *Phys. Rev. Lett.*, **109** (2012) 202501.

339 [26] BRUNO C. G., SCOTT D. A., ALIOTTA M., FORMI-
340 COLA A., BEST A., BOELTZIG A., BEMMERER
341 D., BROGGINI C., CACIOLLI A., CAVANNA F.,
342 CIANI G. F., CORVISIERO P., DAVINSON T., DE-
343 PALO R., DI LEVA A., ELEKES Z., FERRARO F.,
344 FÜLÖP Z., GERVINO G., GUGLIELMETTI A., GUS-
345 TAVINO C., GYÜRKY G., IMBRIANI G., JUNKER M.,
346 MENEGAZZO R., MOSSA V., PANTALEO F. R., PI-
347 ATTI D., PRATI P., SOMORJAI E., STRANIERO O.,
348 STRIEDER F., SZÜCS T., TAKÁCS M. P., TREZZI
349 D. LUNA COLLABORATION, *Phys. Rev. Lett.*, **117**
350 (2016) 142502.